Universo Primitivo
2017-2018 (1º Semestre)

Mestrado em Física - Astronomia

Chapter 1

1. The observed Universe
   - Foreword: The Obler's paradox;
   - The Universe at different scales;
   - Observational Cosmology: empirical facts and the hot Big-Bang theory
     - Cosmic Expansion: The Hubble law;
     - The abundancies of the light elements;
     - The existence of a Cosmic Background Radiation;
     - The isotropy of distant objects;
     - The existent of dark matter;
     - The accelerated expansion of the Universe

Foreword: The Oblers' paradox and our present view about the Universe

Foreword: Why is the sky dark at night?

Heinrich Olbers

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Heinrich Olbers

Olbers' paradox (1826) : argues that "the **darkness of the sky** at night **conflicts with the concept of an eternal and static universe**", with stars distributed uniformly.

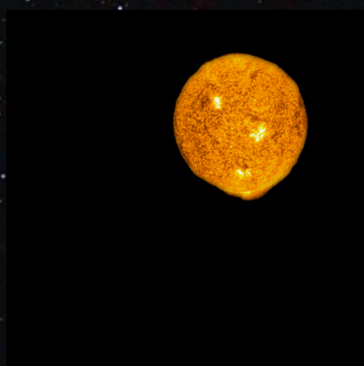
Foreword: Why is the sky dark at night?

Heinrich Olbers

Olbers' paradox (1826) : argues that "the **darkness of the sky** at night **conflicts with the concept of an eternal and static universe**", with stars distributed uniformly.

Oblers paradox in action.

Exercise: prove why this happens



Foreword: Why is the sky dark at night?

Heinrich Olbers

Some possible explanations:

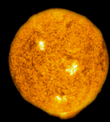
1. Too much dust absorbs light from distant stars.

2. The number of stars in the Universe is finite.

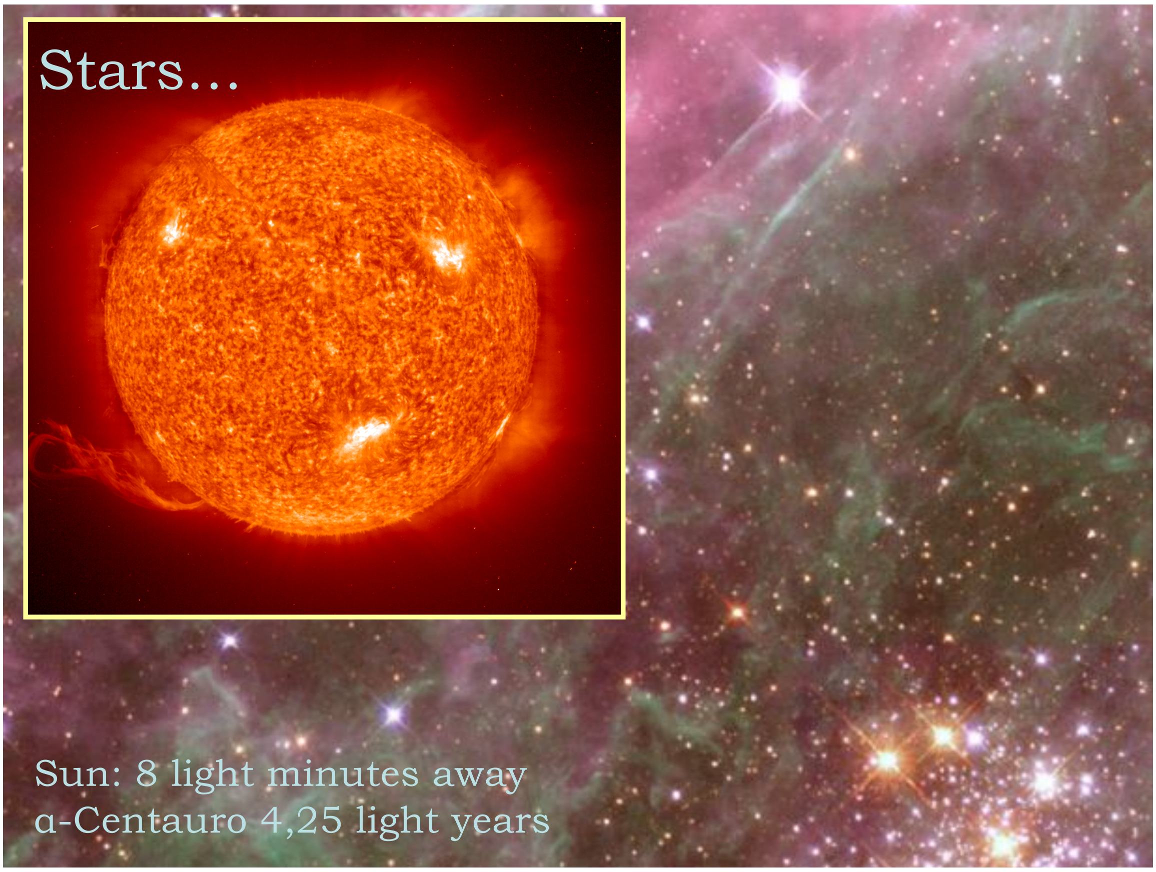
3. The distribution of stars is not uniform.

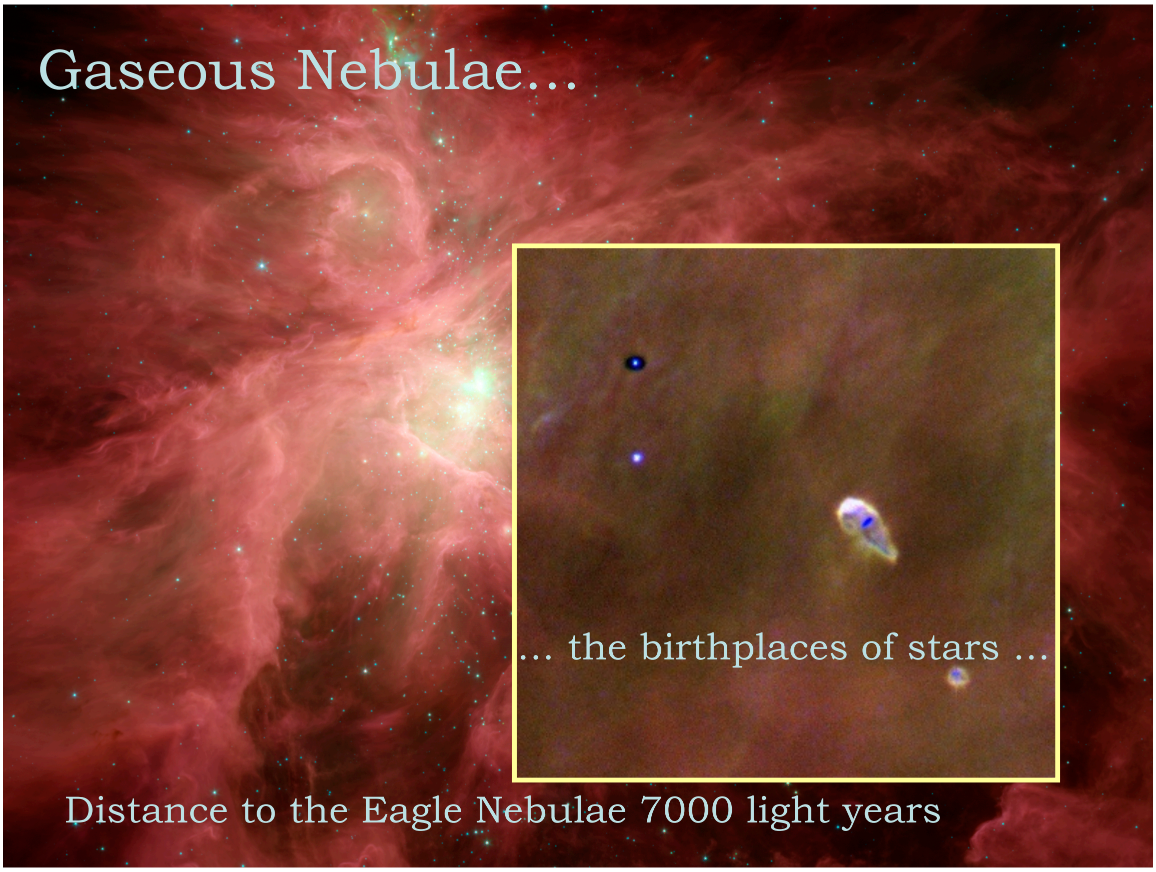
4. The Universe is expanding. Light from distant stars are dimmed (redshifted) into obscurity.

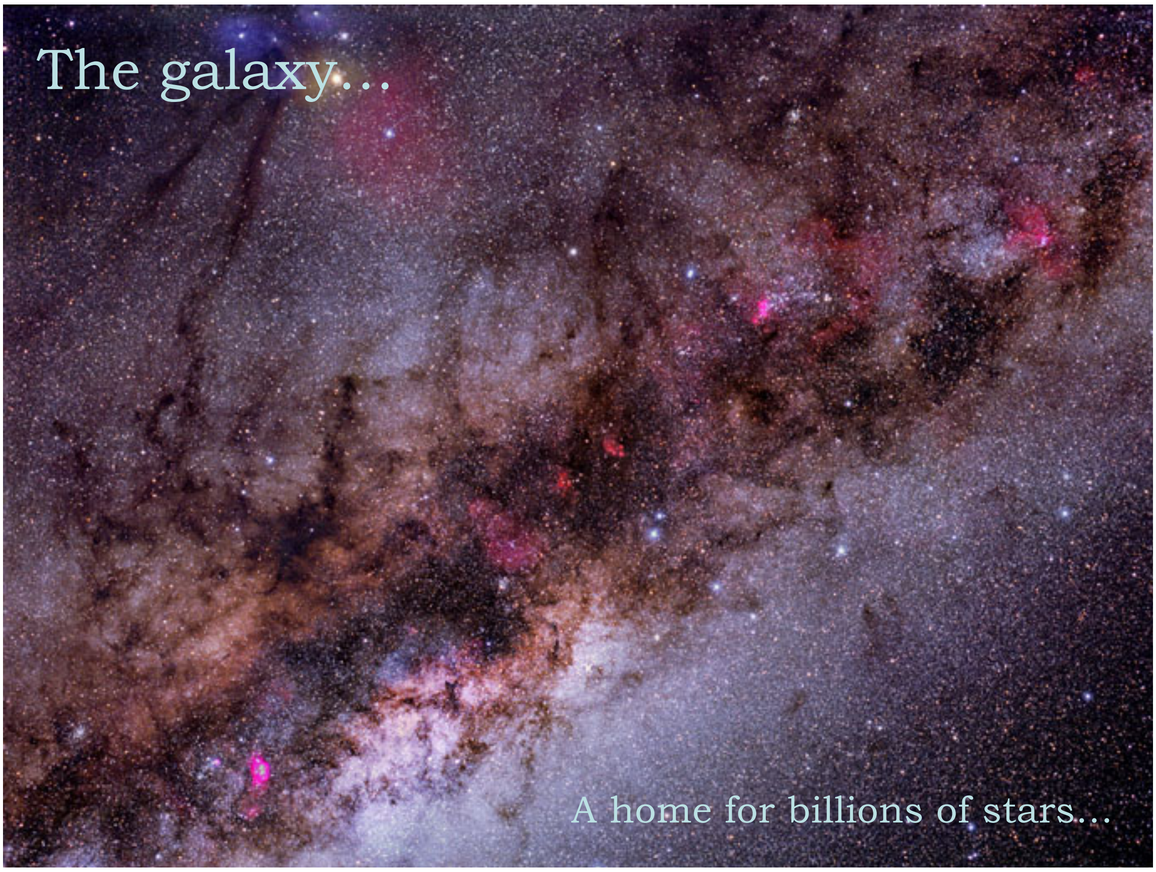
5. The observed Universe has a finite age. Distant light hasn't even reached us yet.

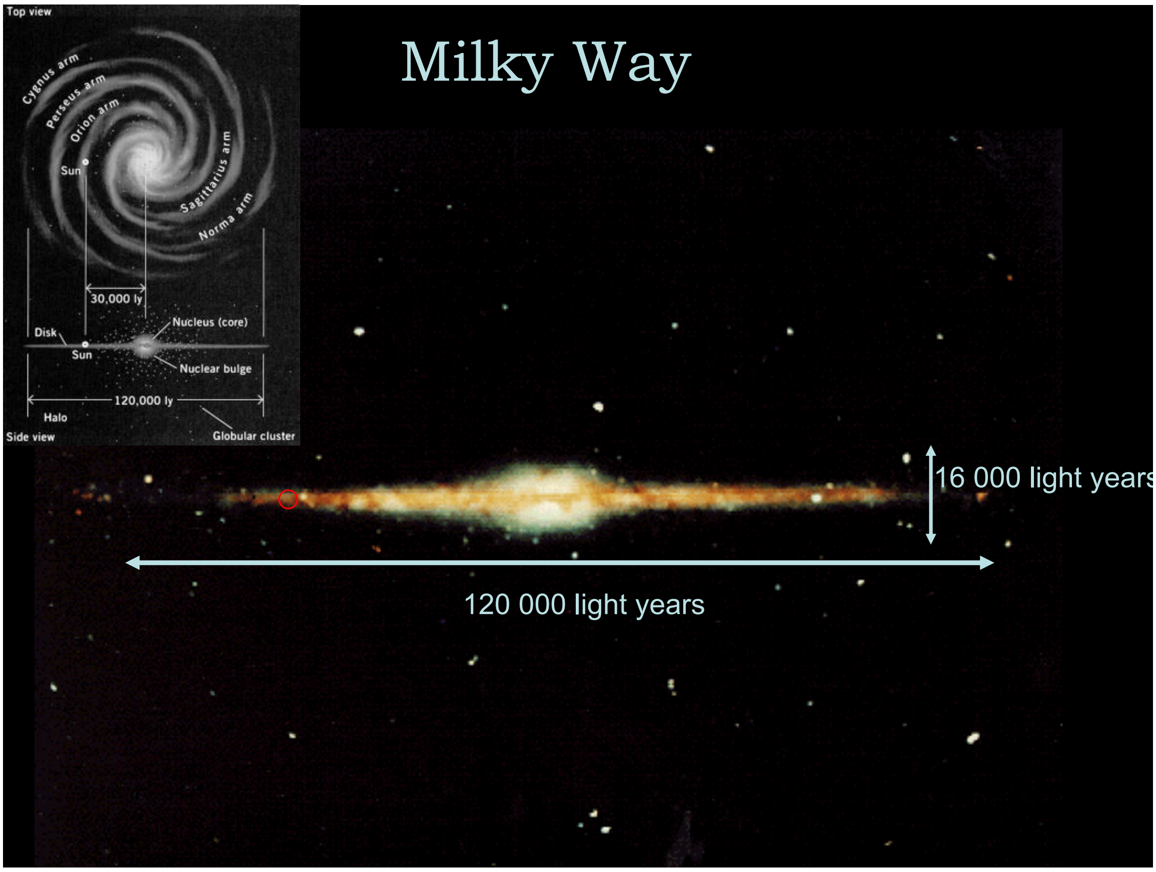


The Universe at different scales

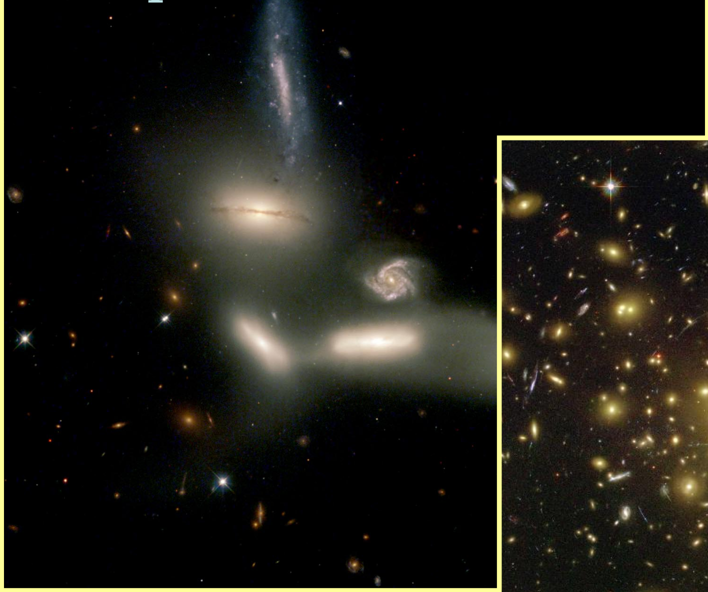

Stars…
Sun: 8 light minutes away
α-Centauro 4,25 light years

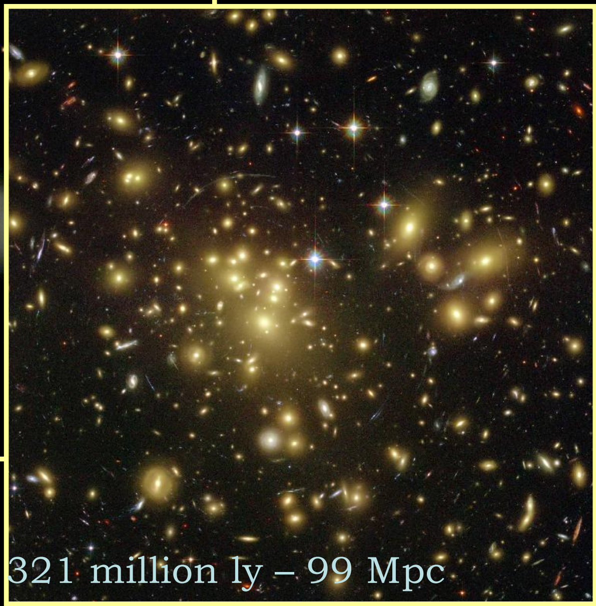

Gaseous Nebulae…
… the birthplaces of stars …
Distance to the Eagle Nebulae 7000 light years


The galaxy…
A home for billions of stars…


Milky Way
Top view
Cygnus arm
Perseus arm
Orion arm
Sun
Sagittarius arm
Norma arm
30,000 ly
Disk
Nucleus (core)
Sun
Nuclear bulge
120,000 ly
Halo
Side view
Globular cluster
16 000 light years
120 000 light years

Groups and Clusters of Galaxies…



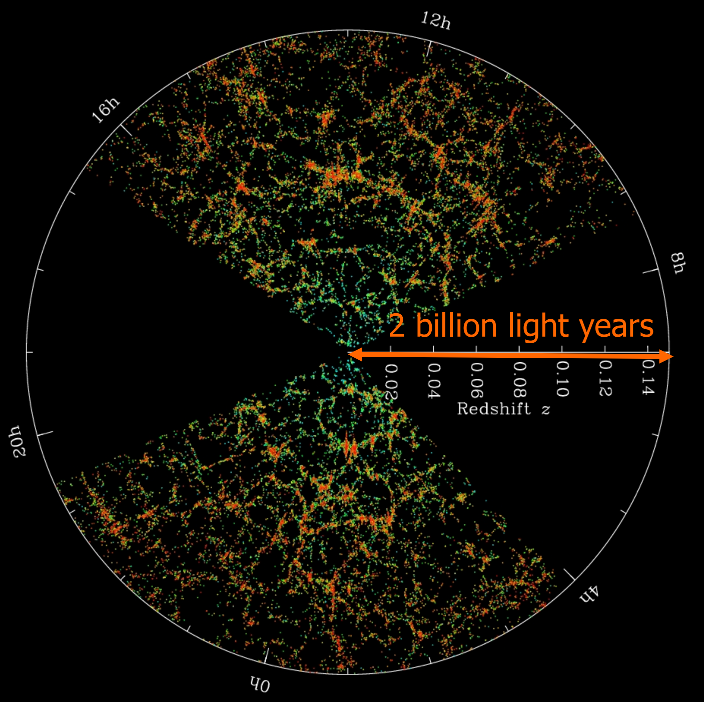


Distance to Coma cluster 321 million ly – 99 Mpc

Our place in the Universe…

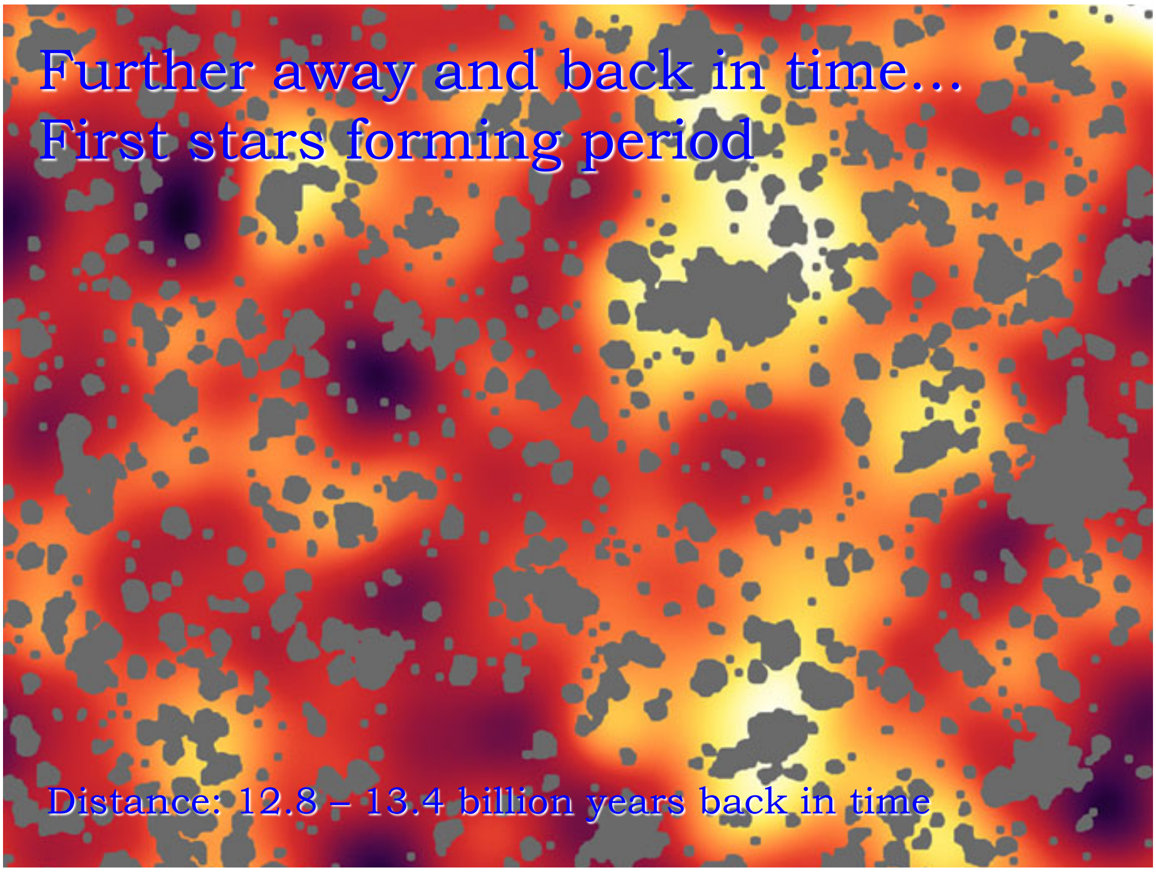
Credits: American Natural History Museum; gently provided by Miguel de Avillez U. Évora

LSS: “The fingers of God”…

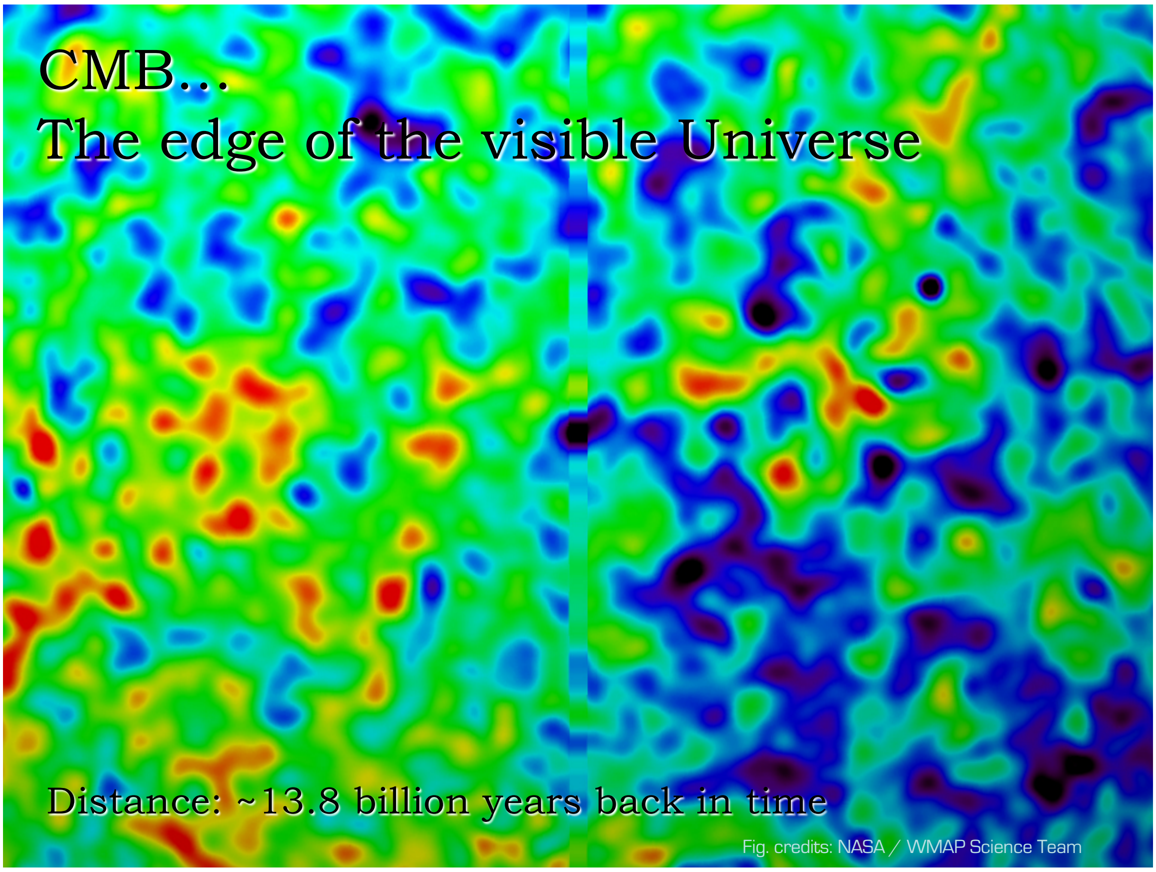


Credit: M. Blanton and the Sloan Digital Sky Survey

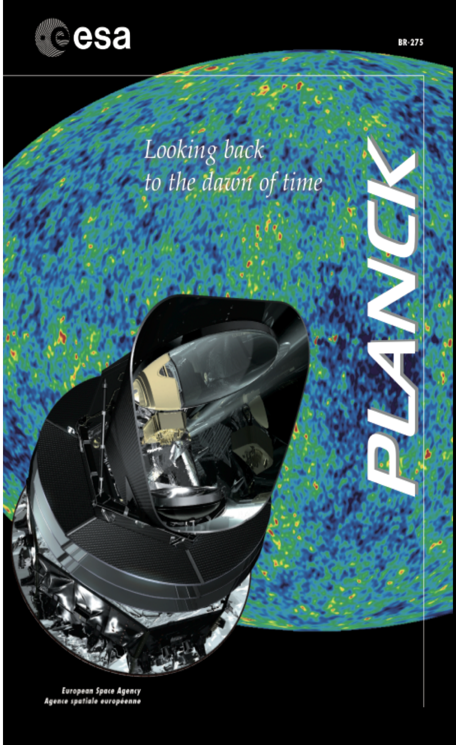
Further away and back in time…
First stars forming period



Distance: 12.8 – 13.4 billion years back in time



Planck Surveyor: looking back to the dawn of time



Project: ESA lead mission to observe the temperature and polarization anisotropies of the Cosmic Microwave Background (CMB) radiation with unprecedented precision.

Total Cost: about €700 million (€1 / person in EU)

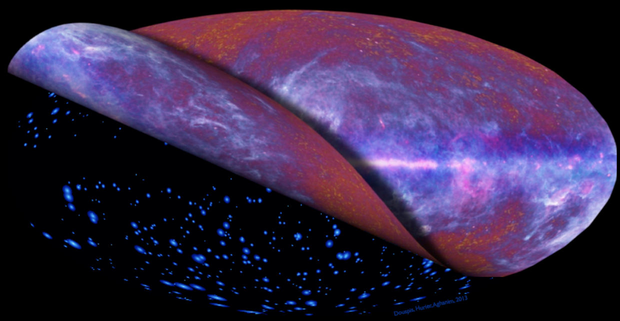
Mission timeline:
- Launch: 14 May 2009
- Operational orbit at L2: July 2009
- Nominal science phase: end of January 2011
- Extended mission: Shut down date: 19 Oct. 2013

Payload:
- Telescope: 1.5 m projected apertures
- Low Frequency Instrument (LFI): array of 22 tuned radio receivers operating at 30, 44 and 70 GHz.
- High Frequency Instrument (HFI): array of 52 bolometers operating at 100, 143, 217, 353, 545, and 857 GHz.

Fig. credits: ESA

Planck CMB observations

2009-2013: Planck satellite observes the CMB sky with unprecedented angular resolution and sensitivity.



Animation credits: ESA and the Planck collaboration; Cluster map by Douspis, Hurier, Aghanim 2013

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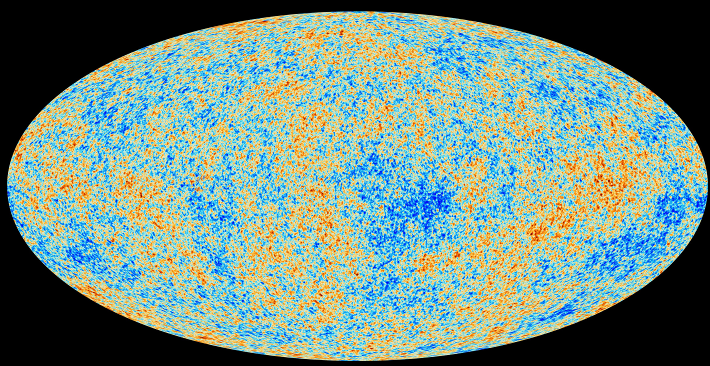
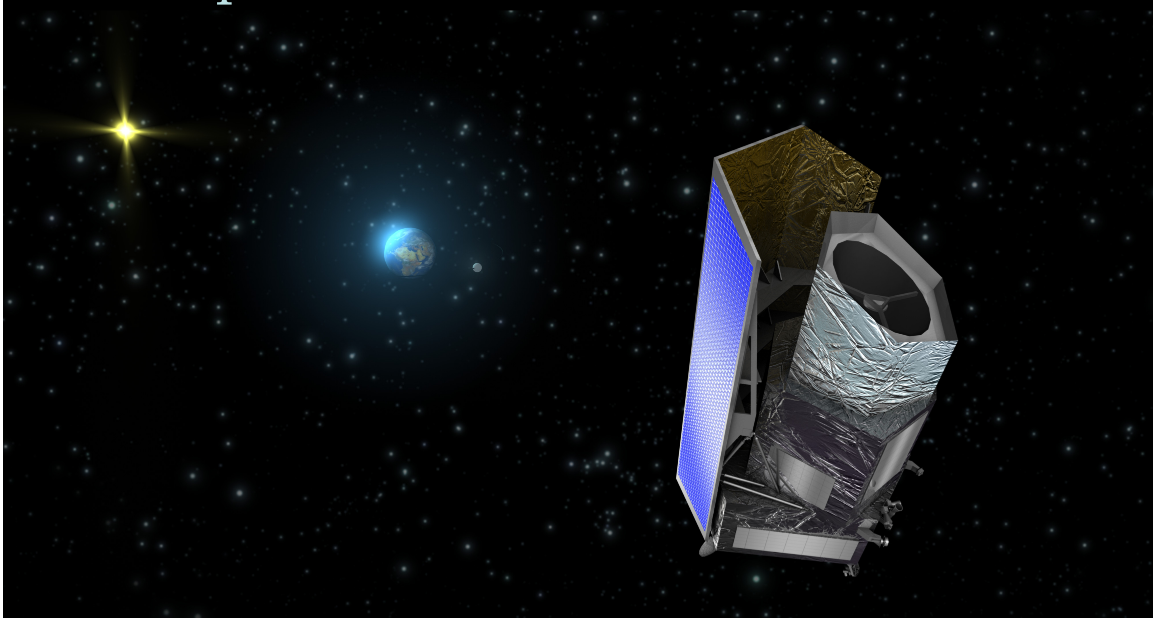


Animation & Fig. credits: ESA and the Planck collaboration

Galaxy surveys: 3D mapping of the Universe…

SDSS: aims at ~25% of the sky; ~100 million objects

Euclid mission (ESA): Galaxy Surveys from space



Euclid: ~ 2000 milions of galaxies

Fig. credits: ESA - C. Carreau.

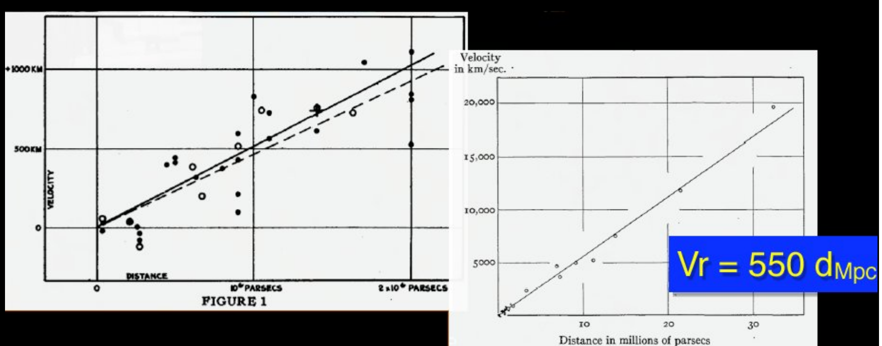
Observational cosmology: empirical facts about the Universe

1. The Universe is expanding

Edwin Hubble

1924: Edwin Hubble ends debate on the nature of nebulae being galactic objects

1929: reports a linear relation between relative radial velocity and distance: $v = Hd$



From:

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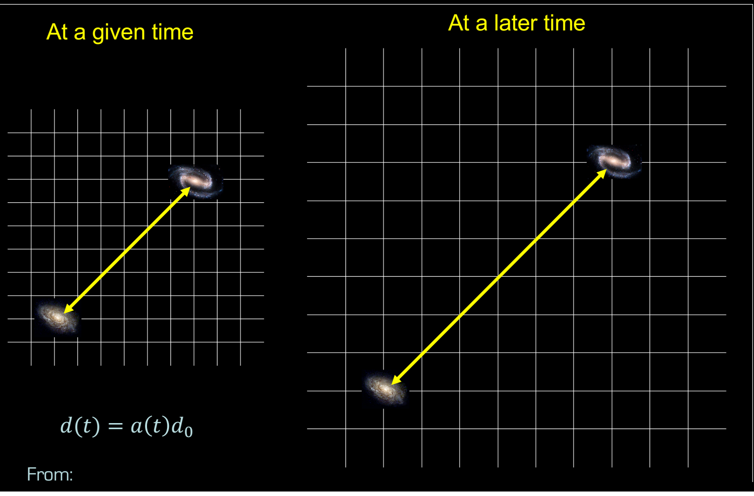

$d(t) = a(t)d_0$

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1. The Universe is expanding

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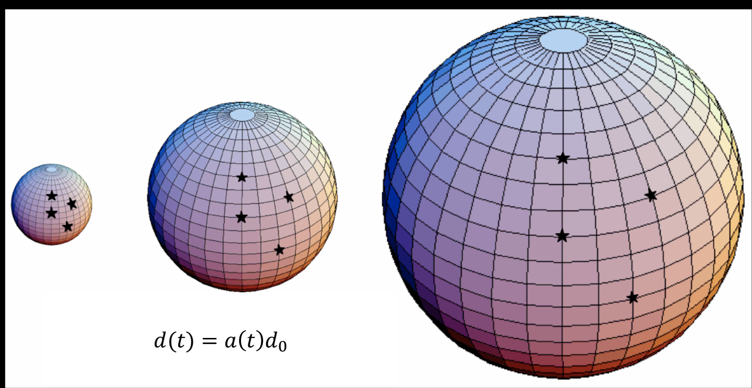
Time evolution of an expanding spherical surface



$d(t) = a(t)d_0$

From:

The basic idea behind the Big-Bang theory

- If the universe is expanding and matter-energy is conserved during the expansion then the universe had to be smaller, denser and hotter in the past!

- If so, the Universe must have evolved from a state where matter and radiation form a ultra dense and hot ionized plasma of fundamental particles

- As the universe expands and cools down:
  - interactions between the plasma components become less frequent;
  - different particle species should decouple from the plasma;
  - eventually the universe becomes neutral and transparent to radiation

According to the Big-Bang theory, in the early instants...

" the Universe was a extremely hot and dense plasma, like a 'torrid bright fog'…

… radiation was trapped in this plasma through collisions with other plasma particles

… as the universe expands, the plasma temperature drops, atomic nuclei form, and capture the free electrons in the plasma. When the number of free electrons is too small, radiation no longer interacts with the plasma and propagates freely, giving rise to the Cosmic Microwave Background and neutral matter"

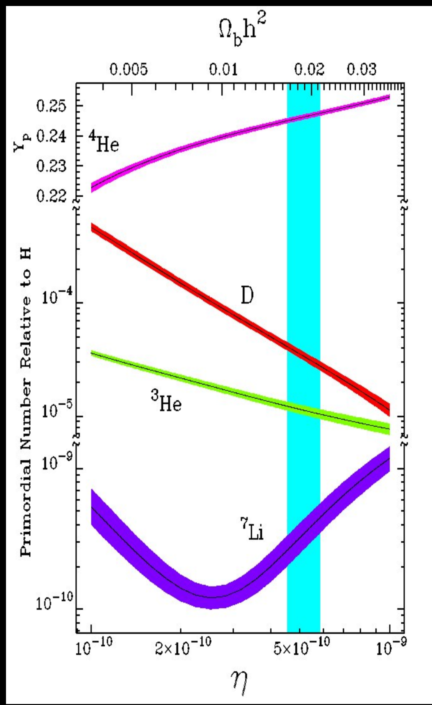
2. The abundance of light nuclei

Herman, Gamow, Alpher

The relative abundance of light elements can not be explained by stellar nucleosyntesis

1948: Alpher & Gamow computed the abundance of light elements in the context of the Big Bang theory

Light elements were produce at low temperatures (<1e9K and high densities) during several tens of minutes



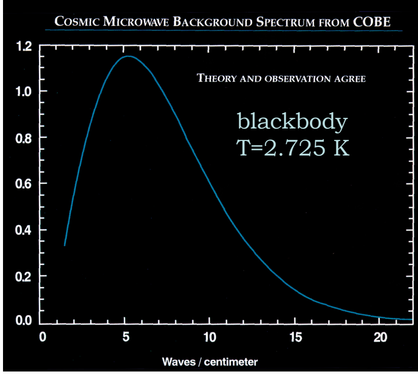
From:

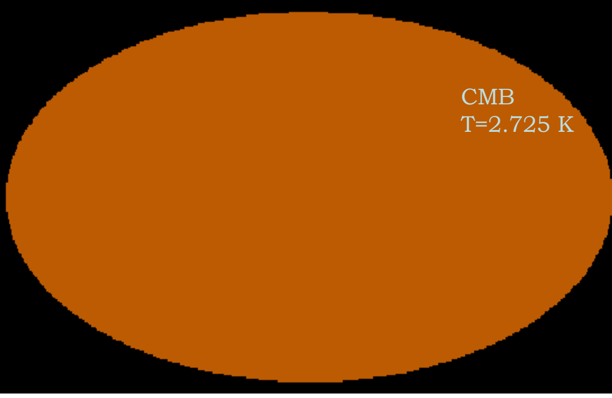
3. Cosmic Microwave Background

Penzias & Wilson

1965: Penzias & Wilson serendipitously discovered a uniform radiation ("excess") across the sky.

This was the cosmic microwave background radiation predicted by Gamow and Alpher in 1948



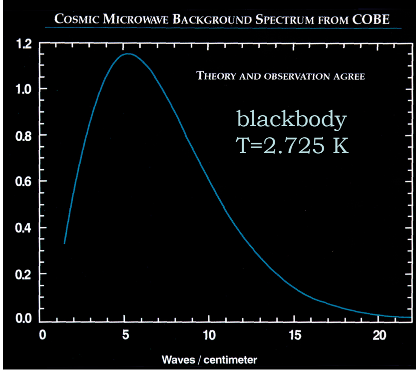


3. Cosmic Microwave Background

John Matter &
Geroge Smooth

1991: High precision measurement of CMB temperature by COBE and 1st detection of temperature fluctuations (Mather & Smoot)

2001: State of the art measurements of dT/T~1e-5 temperature fluctuations by WMAP



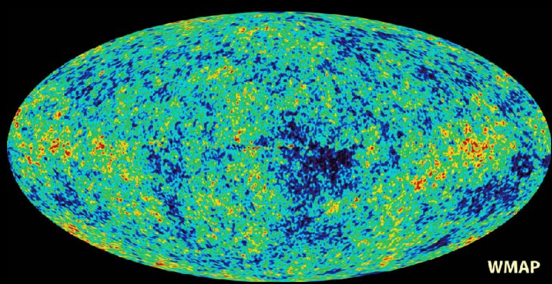
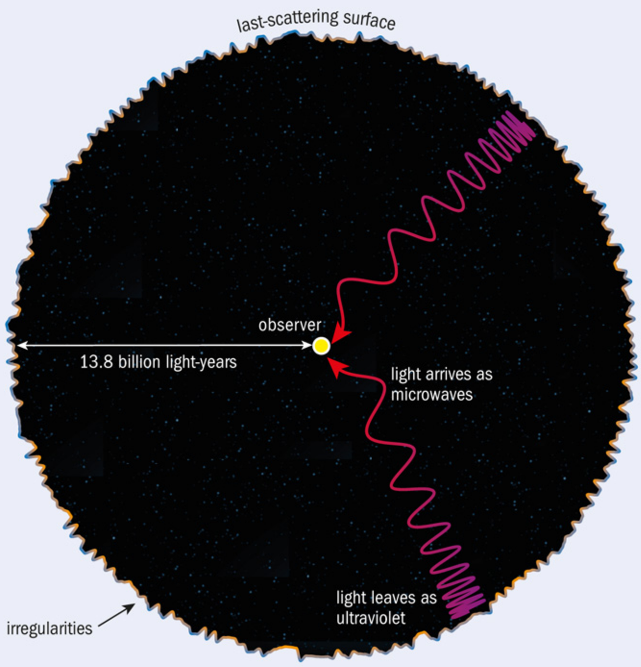
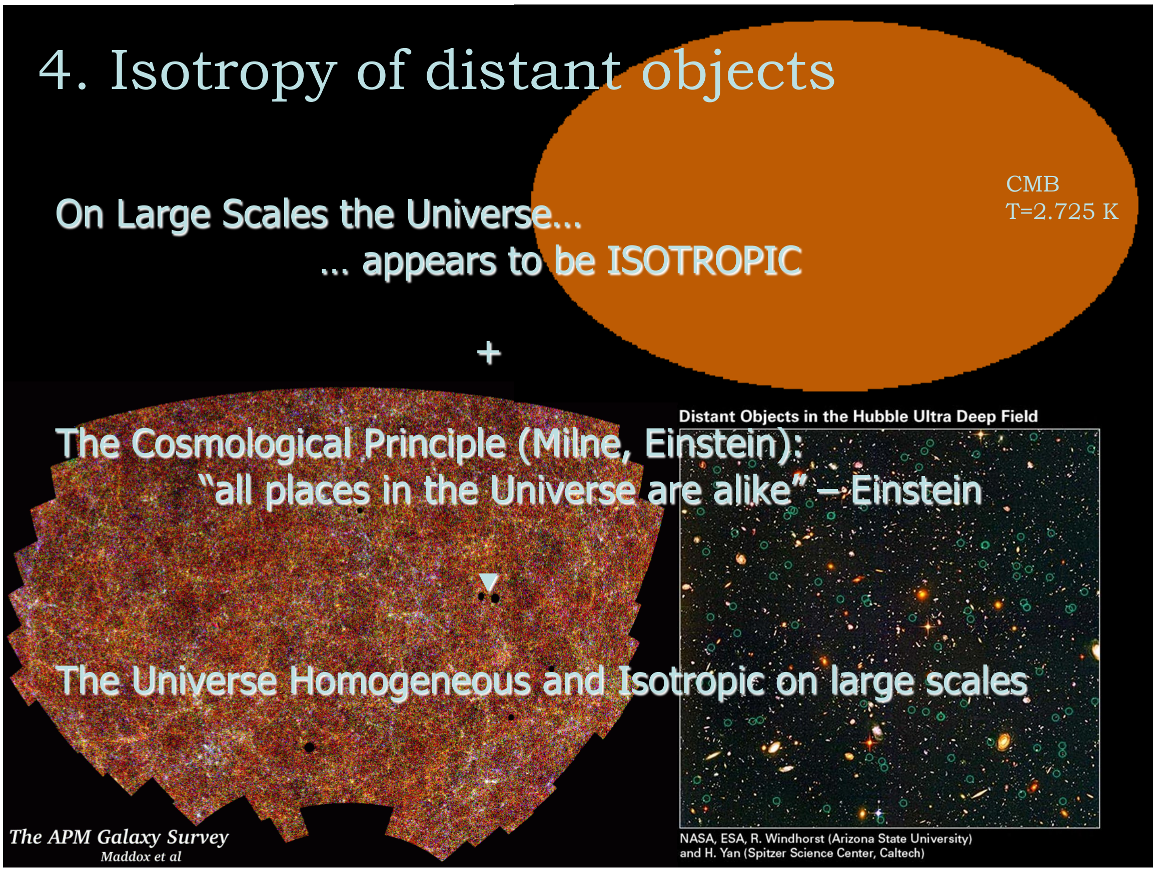


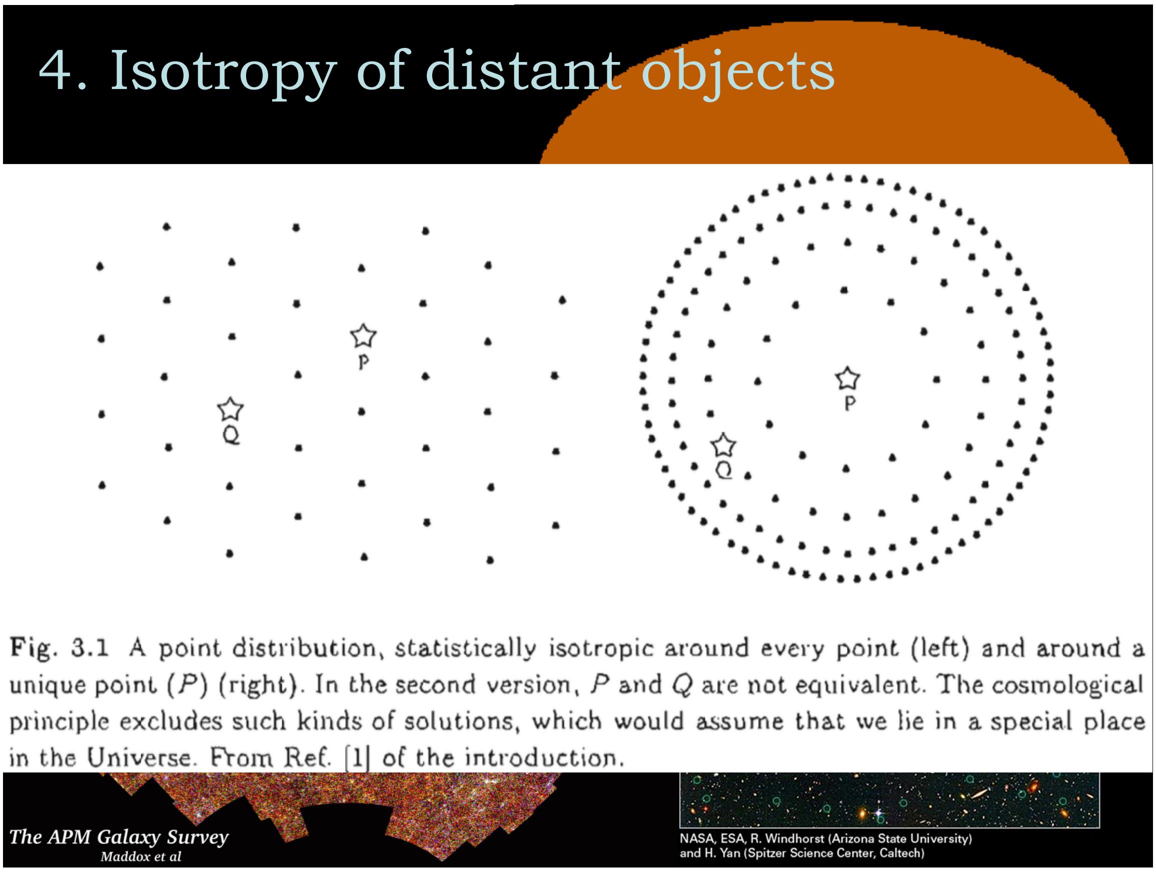
Fig. credits: NASA / WMAP Science Team

CMB: the last scattering surface

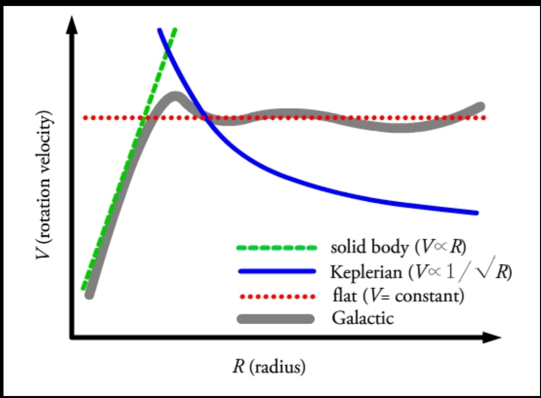


Reprinted from: http://physicsworld.com/cws/article/indepth/2014/jan/09/planck-perspectives


4. Isotropy of distant objects
On Large Scales the Universe...
... appears to be ISOTROPIC
CMB
T=2.725 K
+
Distant Objects in the Hubble Ultra Deep Field
The Cosmological Principle (Milne, Einstein):
"all places in the Universe are alike" – Einstein
The Universe Homogeneous and Isotropic on large scales
The APM Galaxy Survey
Maddox et al
NASA, ESA, R. Windhorst (Arizona State University)
and H. Yan (Spitzer Science Center, Caltech)


4. Isotropy of distant objects
P
Q
P
Q
Fig. 3.1 A point distribution, statistically isotropic around every point (left) and around a unique point (P) (right). In the second version, P and Q are not equivalent. The cosmological principle excludes such kinds of solutions, which would assume that we lie in a special place in the Universe. From Ref. [1] of the introduction.
The APM Galaxy Survey
Maddox et al
NASA, ESA, R. Windhorst (Arizona State University)
and H. Yan (Spitzer Science Center, Caltech)

5. The existence of Dark Matter



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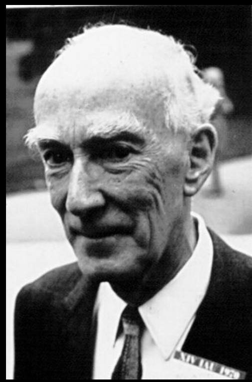


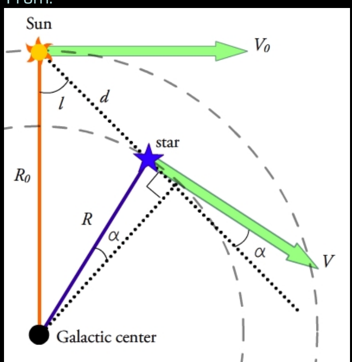
Jan Oort

1927: Jan Oort studies the rotation of stars in our galaxy and infers that their rotation is not consistent with Keplerian motion.

$$v_{circ} = \sqrt{\frac{GM(R)}{r}}$$

5. The existence of Dark Matter

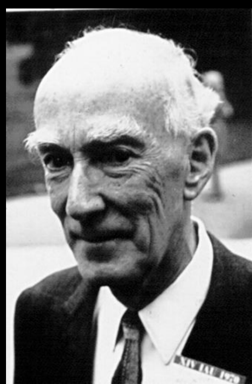
From:



Oorts constants:

$$A \equiv -\frac{1}{2}\left[\frac{dV_c}{dR}|_{R_0} - \frac{V_{c,0}}{R_0}\right]$$

$$B \equiv -\frac{1}{2}\left[\frac{dV_c}{dR}|_{R_0} + \frac{V_{c,0}}{R_0}\right]$$



Circular motion:

$$v_{circ} = \sqrt{\frac{GM(R)}{r}}$$

If the whole mass is mostly at the centre: *vcir ^ 2 ~ 1/r*

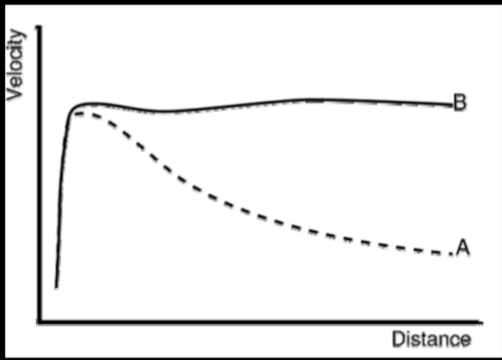
Observations vs Keplerian motion:

- Kepler. motion: (A-B)/(A+B) = 2
- Observations : (A-B)/(A+B) = 5

-Mass is not concentrated at the centre
-Non-luminous mass is required

http://icc.dur.ac.uk/~tt/Lectures/Galaxies/TeX/lec/node42.html

5. The existence of Dark Matter



B: Observations

A: theoretical expectations

1980: Vera Rubin and others also find that stars rotate too fast in the outskirts of spiral galaxies to remain bound assuming that gravity is produced only by visible matter.

5. The existence of Dark Matter

Fritz Zwicky

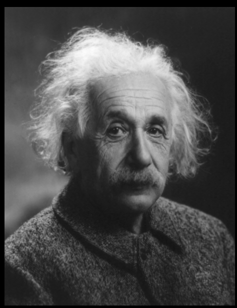
1936: Fritz Zwicky applied the Virial theorem to the velocities of galaxies in the Coma cluster and finds very high Υ =M/L for them to remain bound ($\Upsilon_{coma}/\Upsilon_{sun}$~500 $\gg 2-10$ for galaxies).

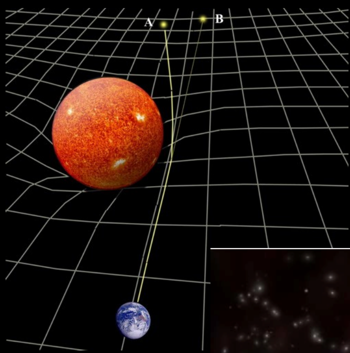
- Viral theorem (for gravitationally relaxed systems): $2\bar{E}_k + \bar{E}_p = 0$
- Mass from the virial theorem: $M_V = < v^2 >< R >/G$
- Visible (luminous) Mass: $M_L = N_g\ R_{ML}\ L_g$

(R_{ML} – typical mass to light ratio of galaxies;
N_g, L_g number and luminosity of individual galaxies)

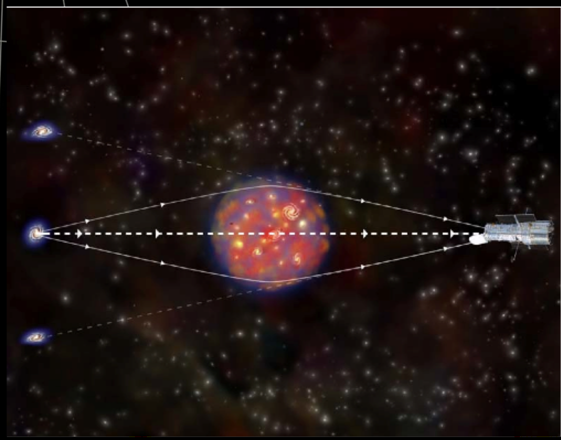
5. The existence of Dark Matter

lensing effects:





$$G_{\mu\nu} = 8\pi G T_{\mu\nu}$$



From:

5. The existence of Dark Matter

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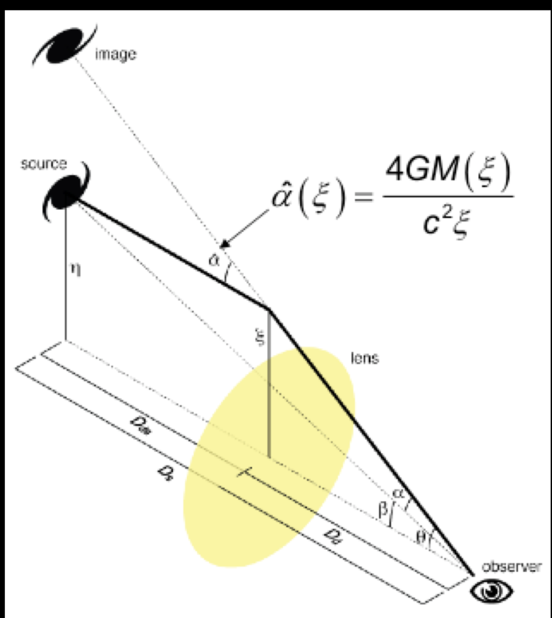


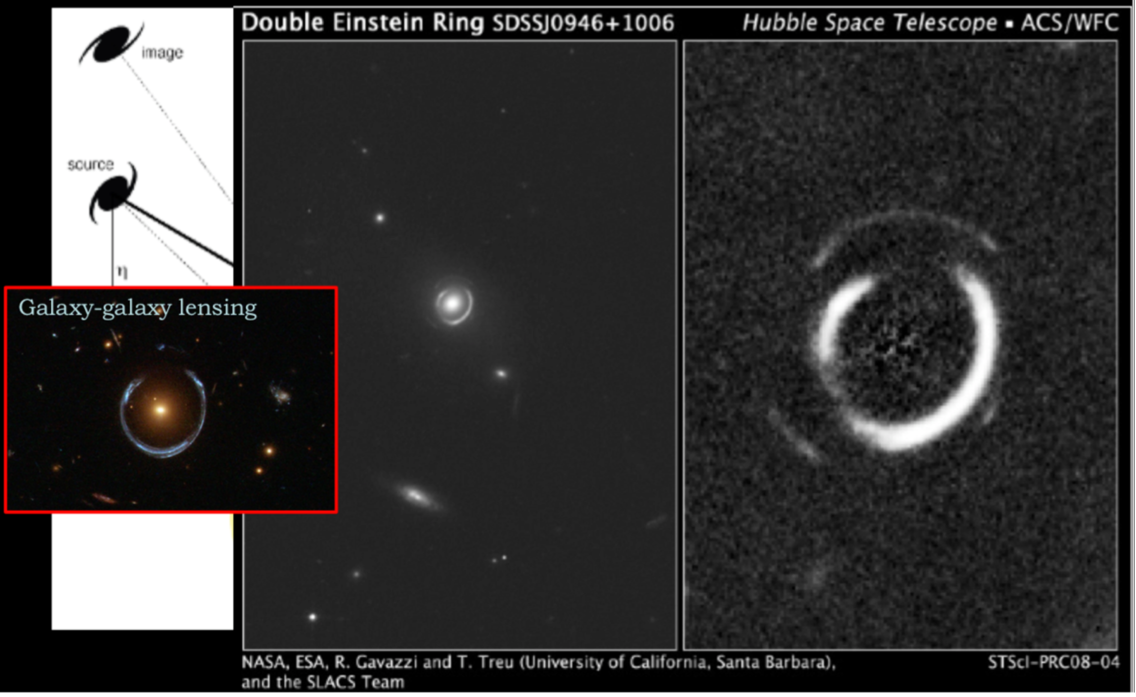
From:

Strong lensing

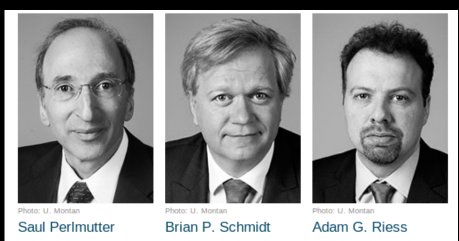
5. The existence of Dark Matter

lensing effects: strong lensing

Einstein Rings



6. Cosmic expansion is accelerating

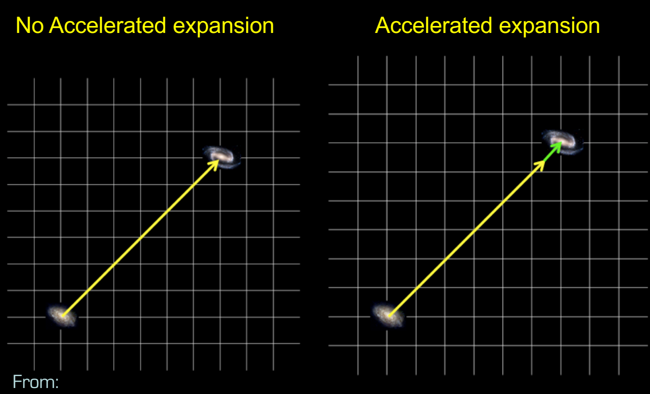


1998: S. Perlmutter and the supernova Cosmology project found first evidence for the accelerated expansion of the Universe.

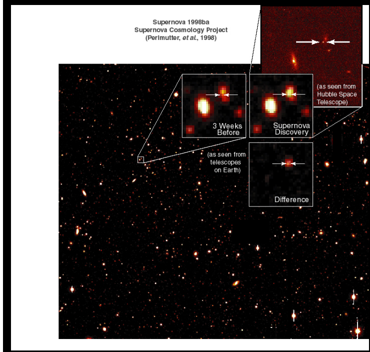
assuming supernovae are standard candles, they appear further away (green arrow) then predicted by non-accelerating expansion models (yellow arrow).

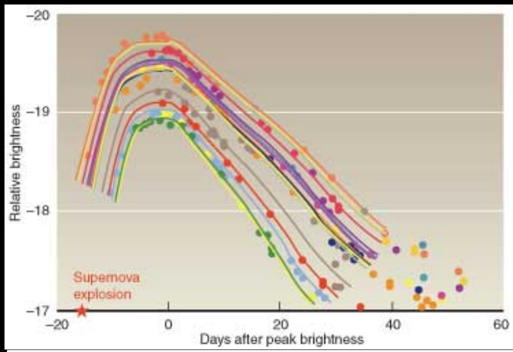
$$d(t) = a(t)\, d_0$$

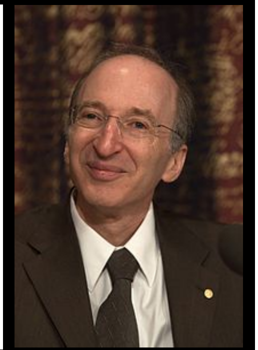
with $\ddot{a}(t) > 0$



6. Cosmic expansion is accelerating

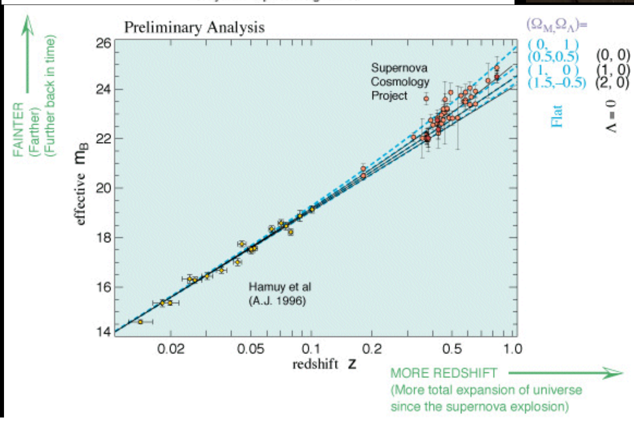




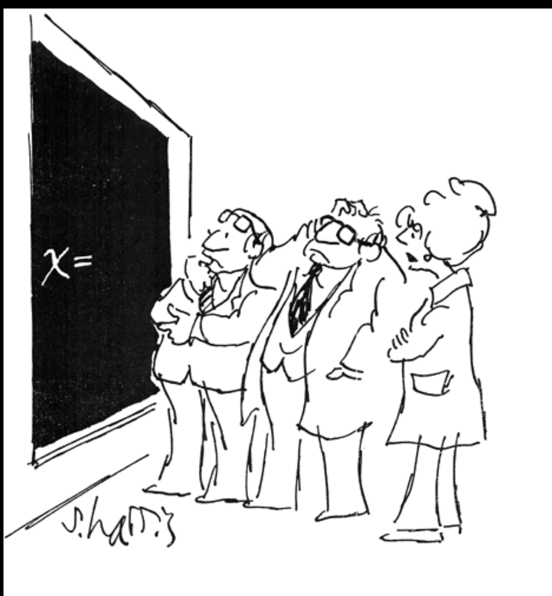


Cosmological redshift:

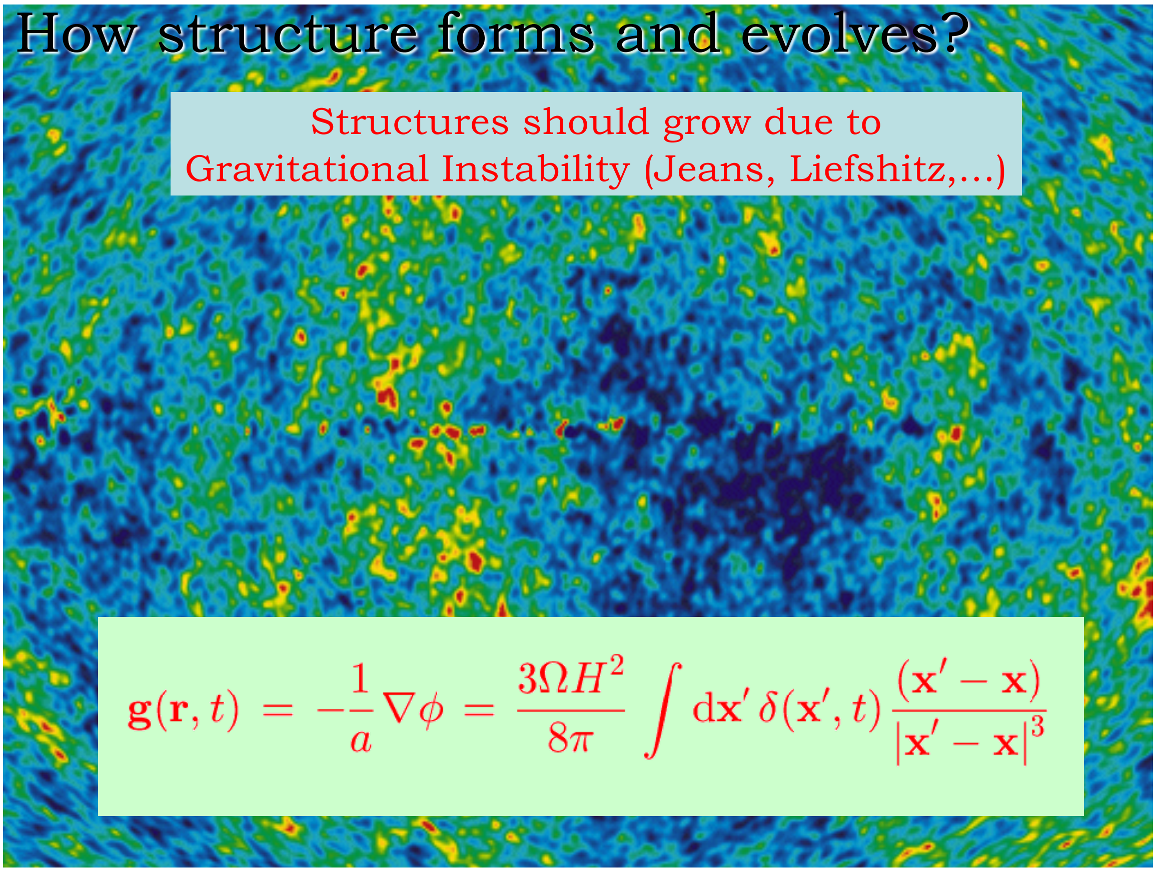
$$z = \frac{E - E_0}{E_0} =$$
$$= \frac{\nu}{\nu_0} - 1 =$$
$$= \frac{\lambda_0}{\lambda} - 1 =$$
$$= \frac{a_0}{a} - 1$$



How Cosmological structure forms and evolves?



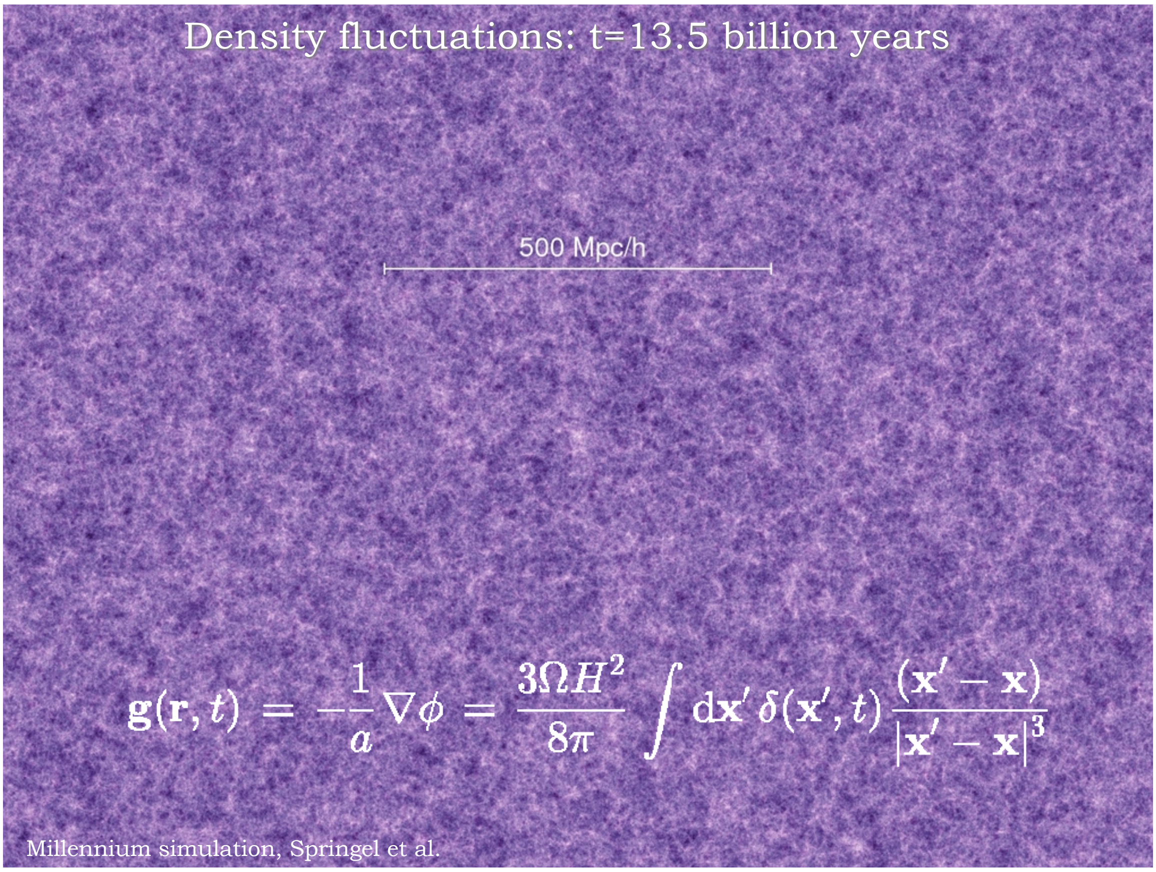
- Observations indicate that
  - on small scales the universe is NOT homogeneous and isotropic
  - On large cosmological scales the Universe doesn't show indications of strong anisotropies. Together with the cosmological principle this implies the universe is highly homogeneous and isotropic
  - However it shows small anisotropies in the CMB.

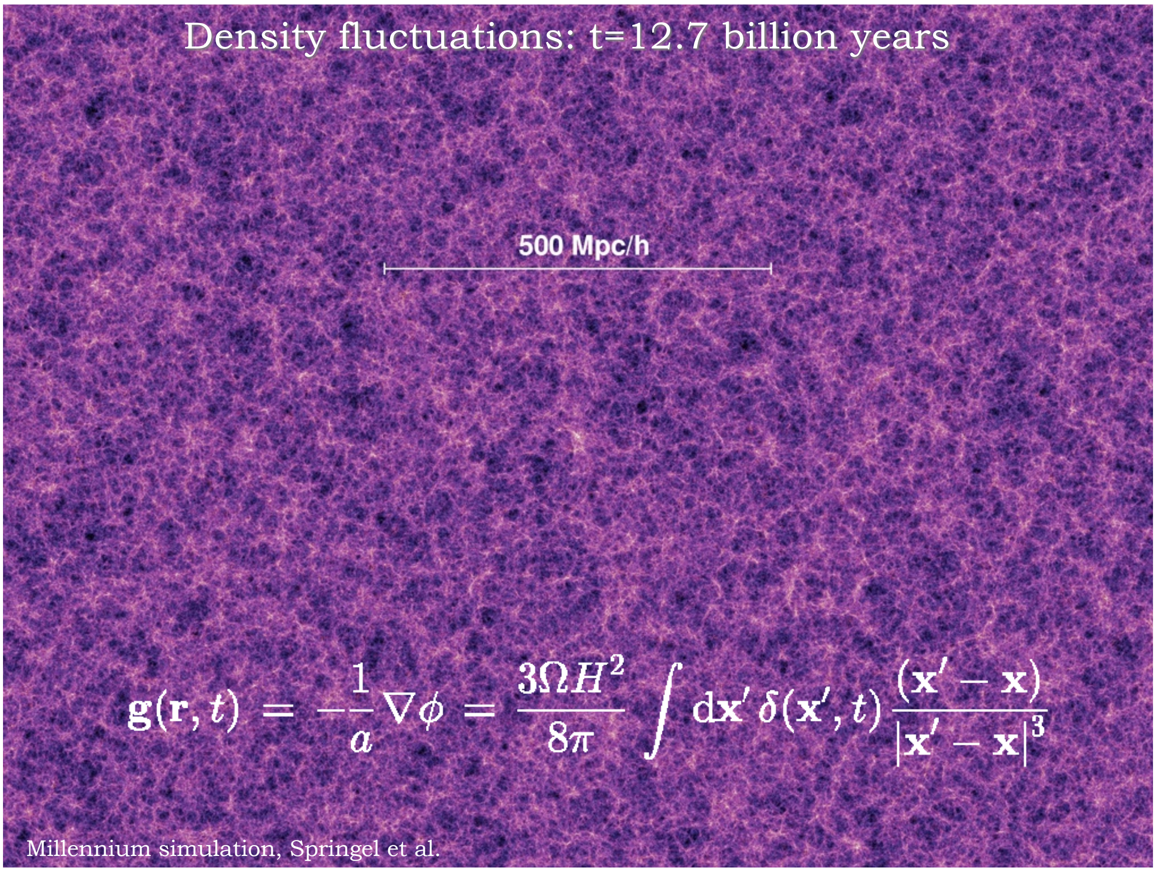


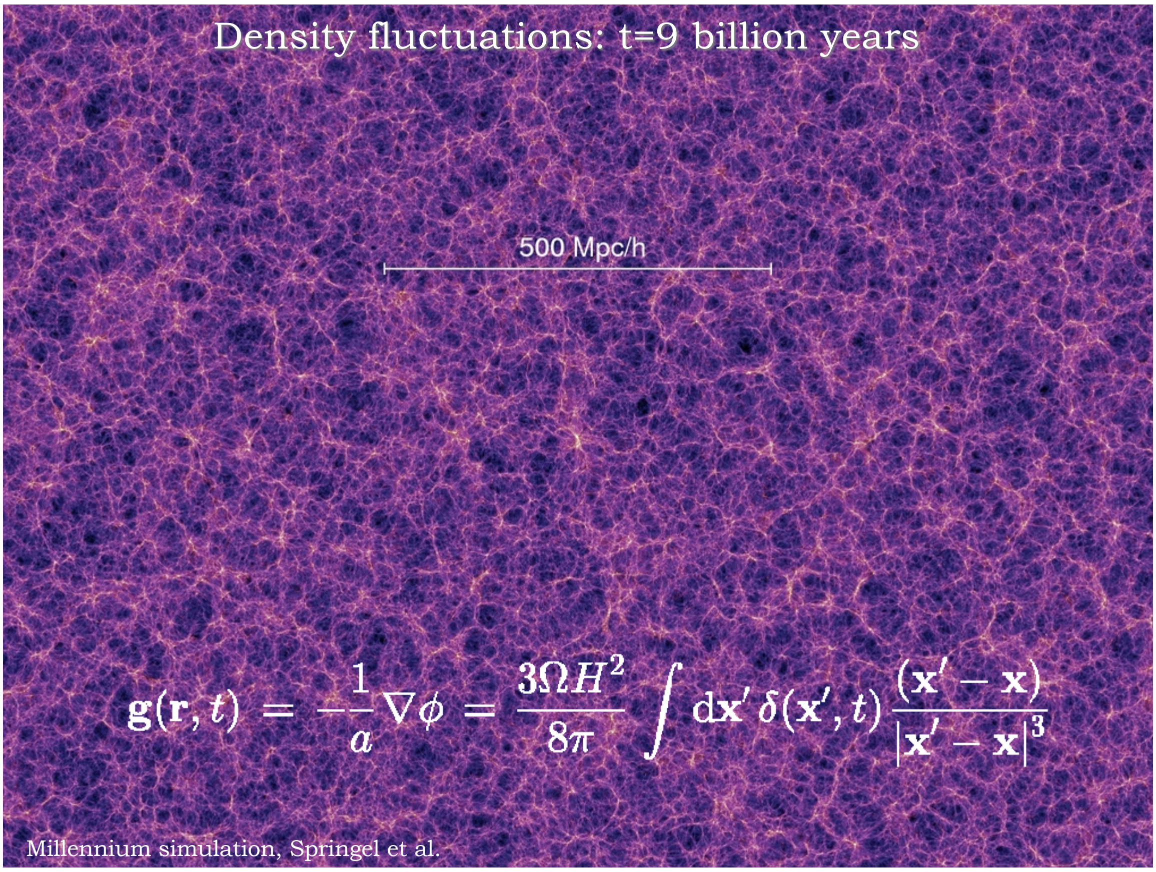
$$\mathbf{g}(\mathbf{r},t) = -\frac{1}{a}\nabla\phi = \frac{3\Omega H^2}{8\pi}\int \mathrm{d}\mathbf{x}'\,\delta(\mathbf{x}',t)\,\frac{(\mathbf{x}'-\mathbf{x})}{|\mathbf{x}'-\mathbf{x}|^3}$$

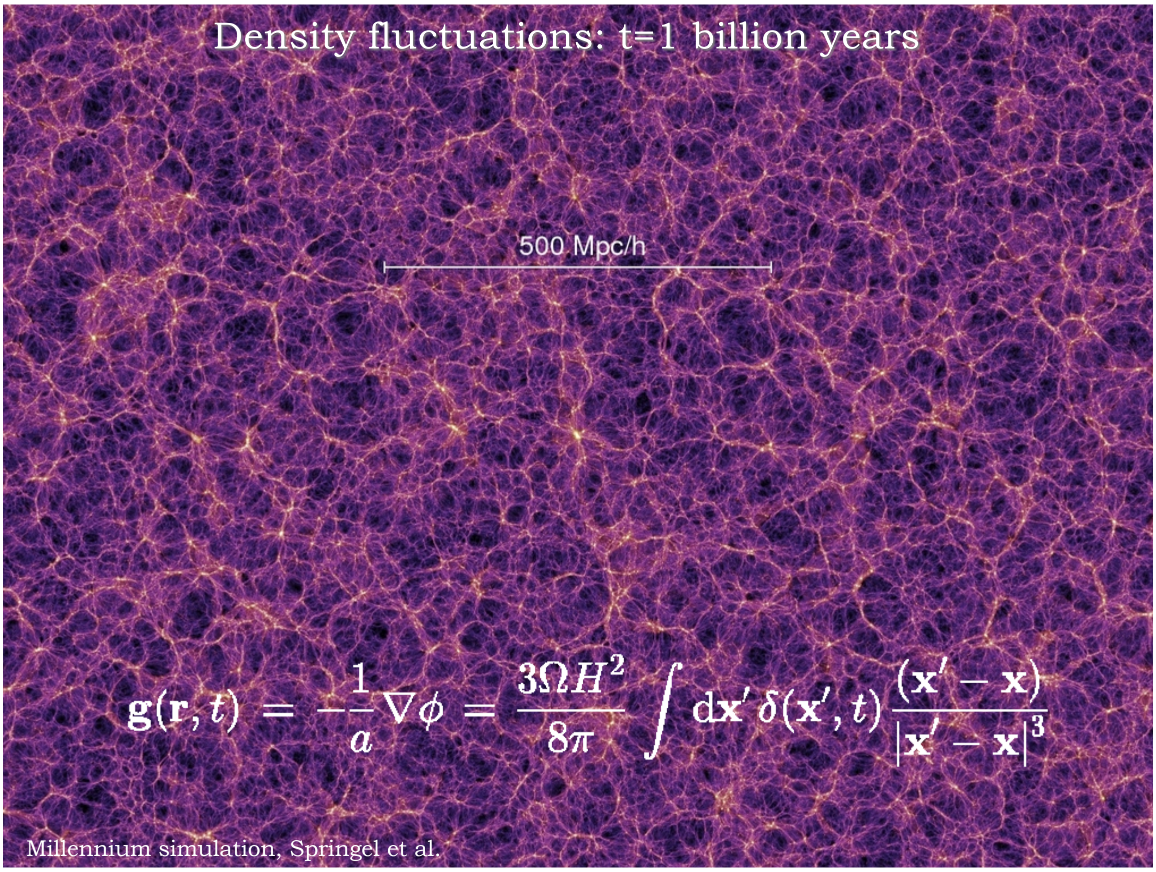
Density fluctuations: t=13.7 billion years

$$\mathbf{g}(\mathbf{r},t) = -\frac{1}{a}\nabla\phi = \frac{3\Omega H^2}{8\pi}\int \mathrm{d}\mathbf{x}'\,\delta(\mathbf{x}',t)\,\frac{(\mathbf{x}'-\mathbf{x})}{|\mathbf{x}'-\mathbf{x}|^3}$$


Density fluctuations: t=13.5 billion years
500 Mpc/h
$$\mathbf{g}(\mathbf{r},t) = -\frac{1}{a}\nabla\phi = \frac{3\Omega H^2}{8\pi}\int \mathrm{d}\mathbf{x}'\,\delta(\mathbf{x}',t)\,\frac{(\mathbf{x}'-\mathbf{x})}{|\mathbf{x}'-\mathbf{x}|^3}$$
Millennium simulation, Springel et al.

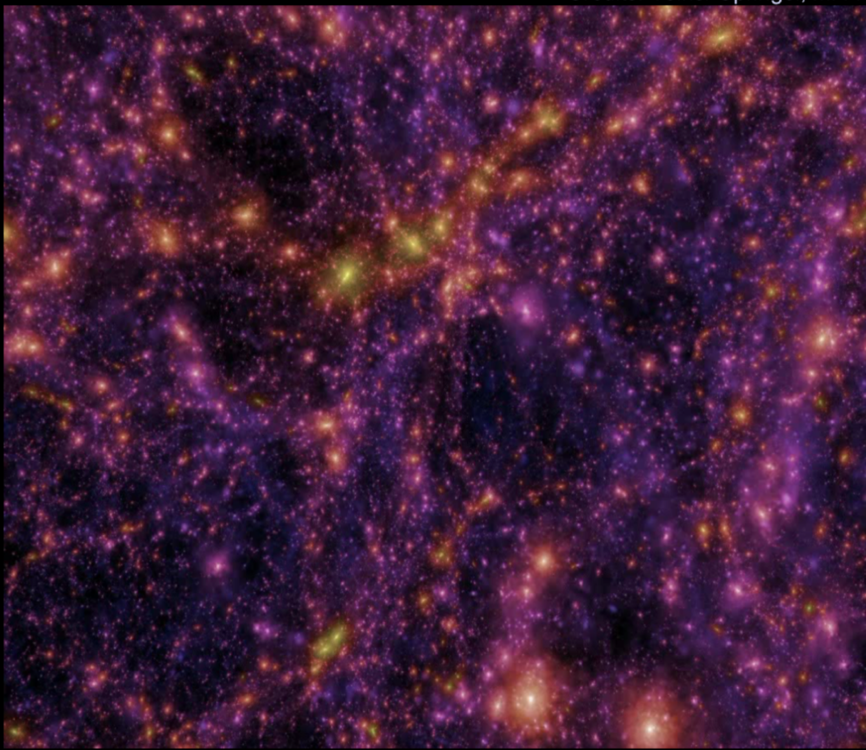

Density fluctuations: t=12.7 billion years
500 Mpc/h
$$\mathbf{g}(\mathbf{r},t) = -\frac{1}{a}\nabla\phi = \frac{3\Omega H^2}{8\pi}\int \mathrm{d}\mathbf{x}'\,\delta(\mathbf{x}',t)\,\frac{(\mathbf{x}'-\mathbf{x})}{|\mathbf{x}'-\mathbf{x}|^3}$$
Millennium simulation, Springel et al.


Density fluctuations: t=9 billion years
500 Mpc/h
$$\mathbf{g}(\mathbf{r},t) = -\frac{1}{a}\nabla\phi = \frac{3\Omega H^2}{8\pi}\int \mathrm{d}\mathbf{x}'\,\delta(\mathbf{x}',t)\,\frac{(\mathbf{x}'-\mathbf{x})}{|\mathbf{x}'-\mathbf{x}|^3}$$
Millennium simulation, Springel et al.


Density fluctuations: t=1 billion years
500 Mpc/h
$$\mathbf{g}(\mathbf{r},t) = -\frac{1}{a}\nabla\phi = \frac{3\Omega H^2}{8\pi}\int \mathrm{d}\mathbf{x}'\,\delta(\mathbf{x}',t)\,\frac{(\mathbf{x}'-\mathbf{x})}{|\mathbf{x}'-\mathbf{x}|^3}$$
Millennium simulation, Springel et al.

Density fluctuations: local universe

Credits: Volker Springel, MPA



The history of the Universe: