Universo Primitivo

2022-2023 (1º Semestre)

Mestrado em Física - Astronomia

Chapter 1

1. The observed Universe
   - Foreword: The Olbers' paradox;
   - The Universe at different scales;
   - Observational Cosmology: empirical facts and the hot Big-Bang theory
     - Cosmic Expansion: The Hubble law;
     - The abundancies of the light elements;
     - The existence of a Cosmic Background Radiation;
     - The isotropy of distant objects;
     - The existent of dark matter;
     - The accelerated expansion of the Universe
   - Formation and evolution of cosmic structure

Foreword: The Olbers' paradox and the present view of the Universe

Foreword: Why is the sky dark at night?

Heinrich Olbers
(1758–1840)

Foreword: Why is the sky dark at night?

Heinrich Olbers

(1758–1840

Olbers' paradox (1826) : argues that "the **darkness of the sky** at night **conflicts with the concept of an infinite and eternal static universe**", with stars distributed uniformly.

$$\text{light} = \int_{r_0}^{\infty} L(r)N(r)\,dr$$

Foreword: Why is the sky dark at night?

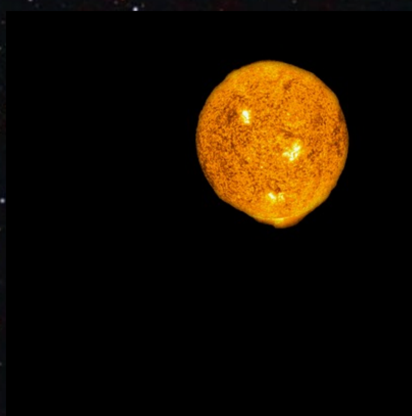
Heinrich Olbers

(1758–1840

Olbers' paradox (1826) : argues that "the **darkness of the sky** at night **conflicts with the concept of an infinite and eternal static universe**", with stars distributed uniformly.

Oblers paradox in action.

Exercise: prove why this happens



Foreword: Why is the sky dark at night?

Heinrich Olber's

(1758–1840

Some possible explanations:

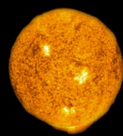
1. Too much dust absorbs light from distant stars.

2. The number of stars in the Universe is finite.

3. The distribution of stars is not uniform.

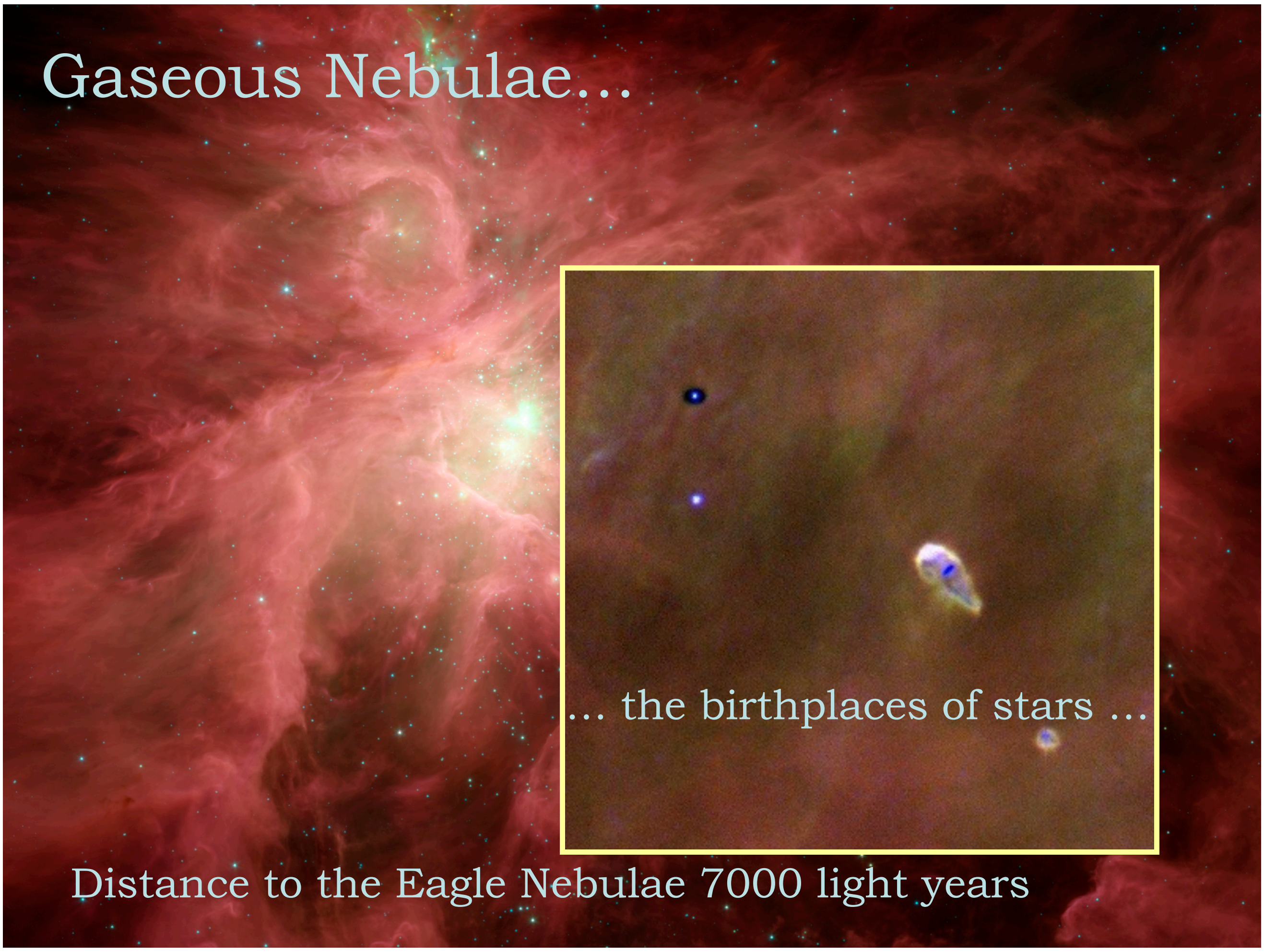
4. The Universe is expanding. Light from distant stars are dimmed (redshifted) into obscurity.

5. The observed Universe has a finite age. Distant light hasn't even reached us yet.

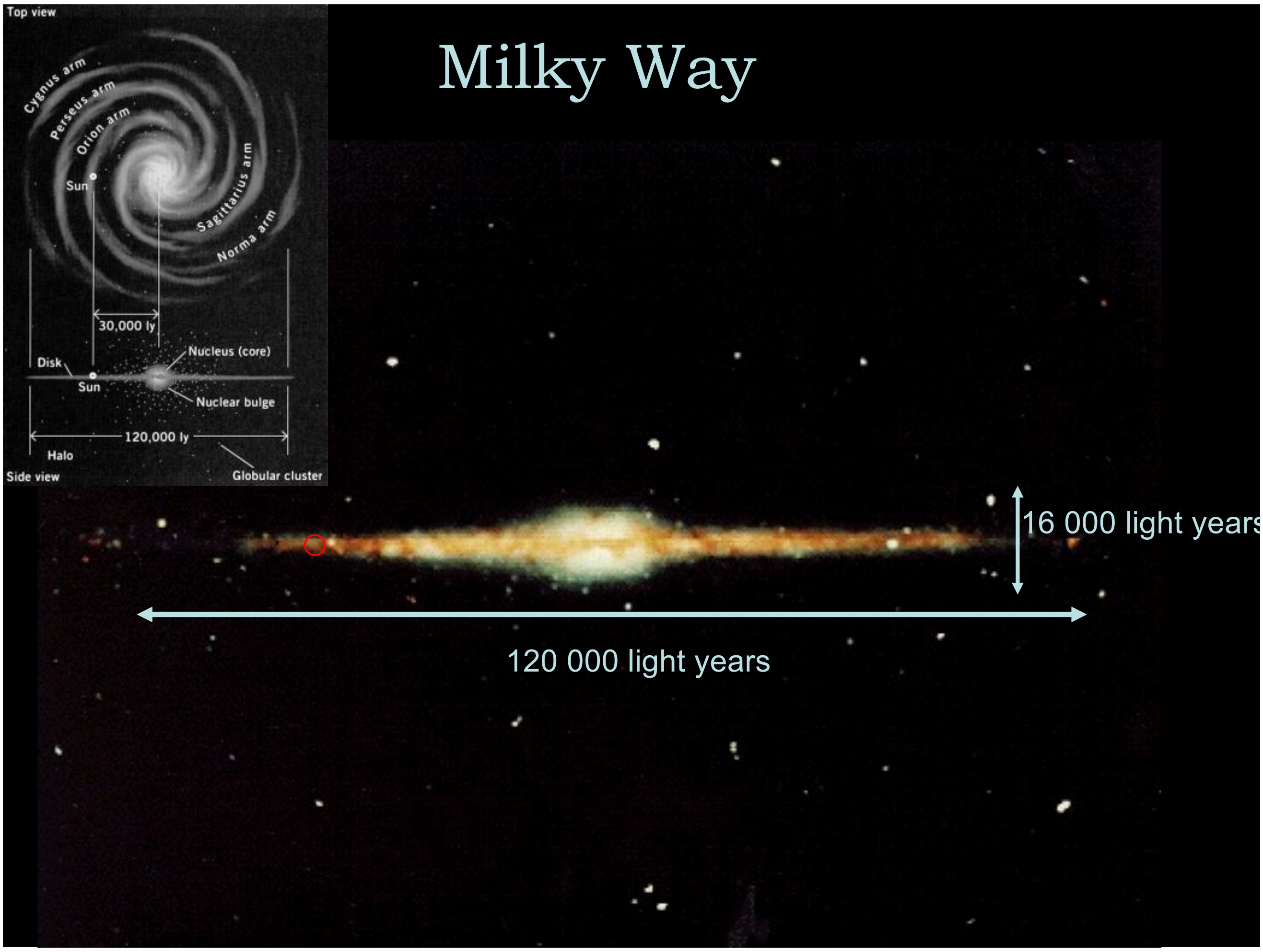


The Universe at different scales

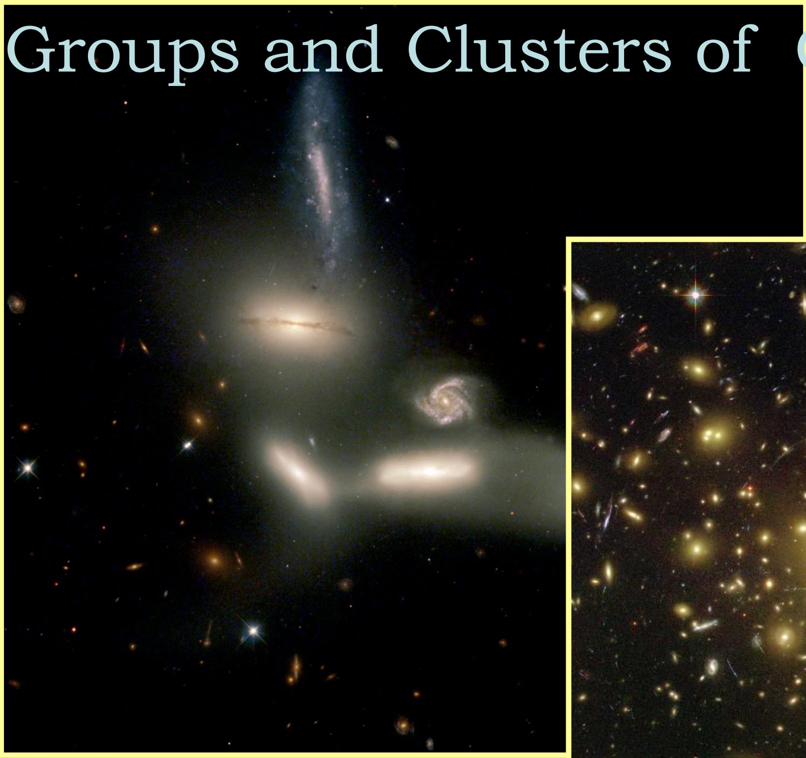

Stars…
Sun: 8 light minutes away
α-Centauro 4,25 light years

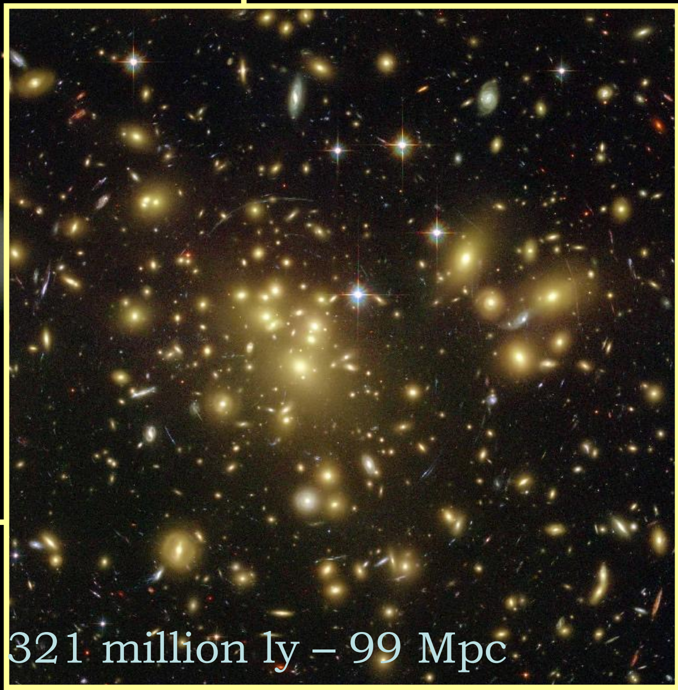

Gaseous Nebulae…
… the birthplaces of stars …
Distance to the Eagle Nebulae 7000 light years


The galaxy…
A home for billions of stars…


Milky Way
Top view
Cygnus arm
Perseus arm
Orion arm
Sun
Sagittarius arm
Norma arm
30,000 ly
Disk
Nucleus (core)
Sun
Nuclear bulge
120,000 ly
Halo
Side view
Globular cluster
16 000 light years
120 000 light years

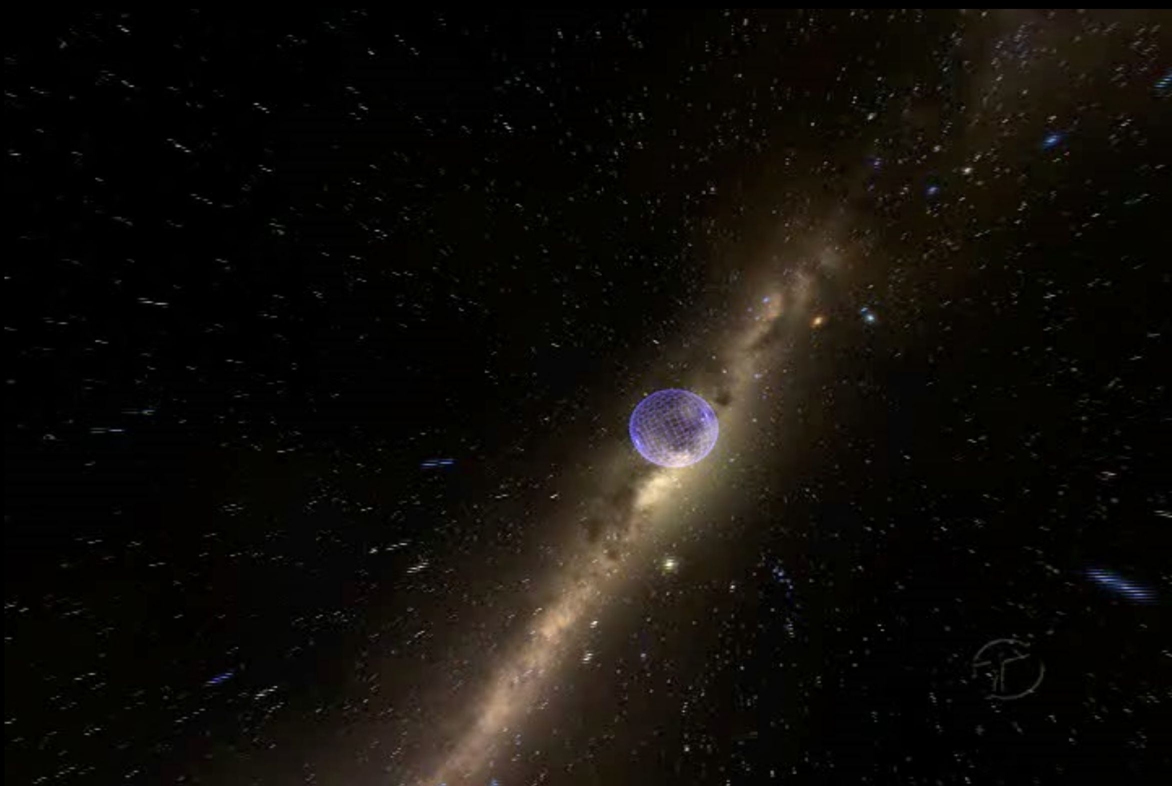
Groups and Clusters of Galaxies…





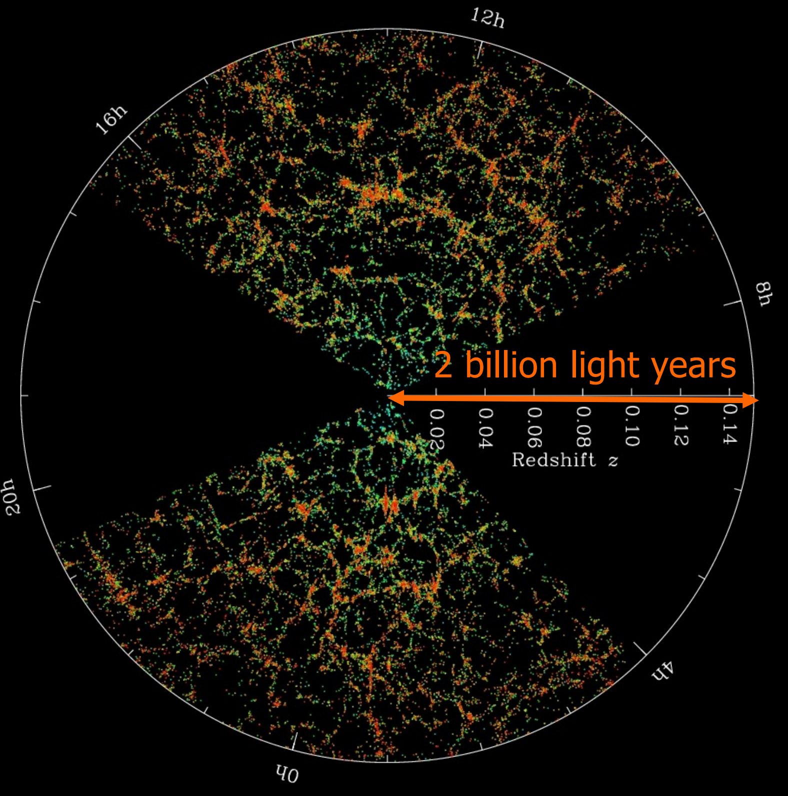
Distance to Coma cluster 321 million ly – 99 Mpc

Our place in the Universe…



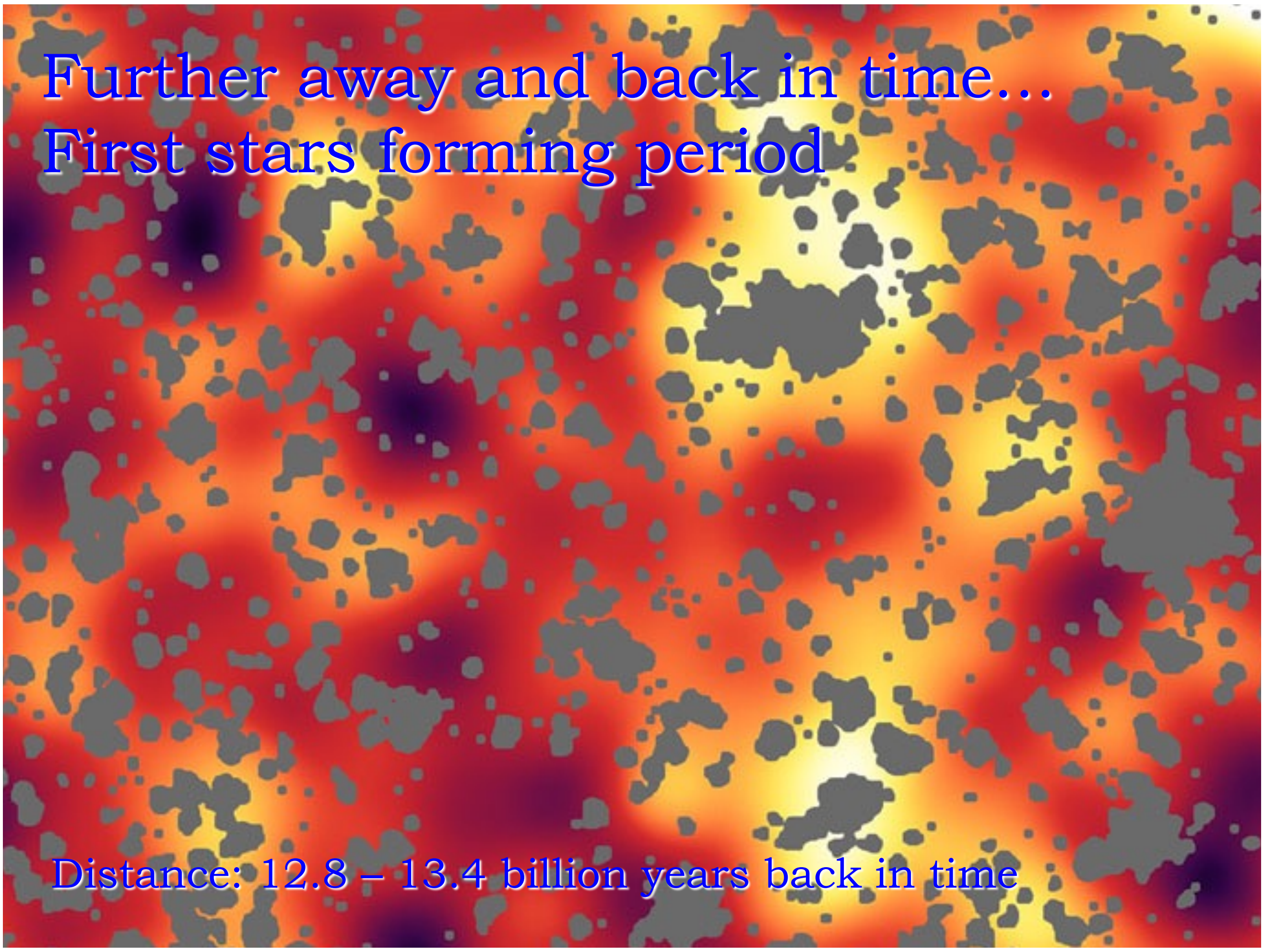
Credits: American Natural History Museum; gently provided by Miguel de Avillez U. Évora

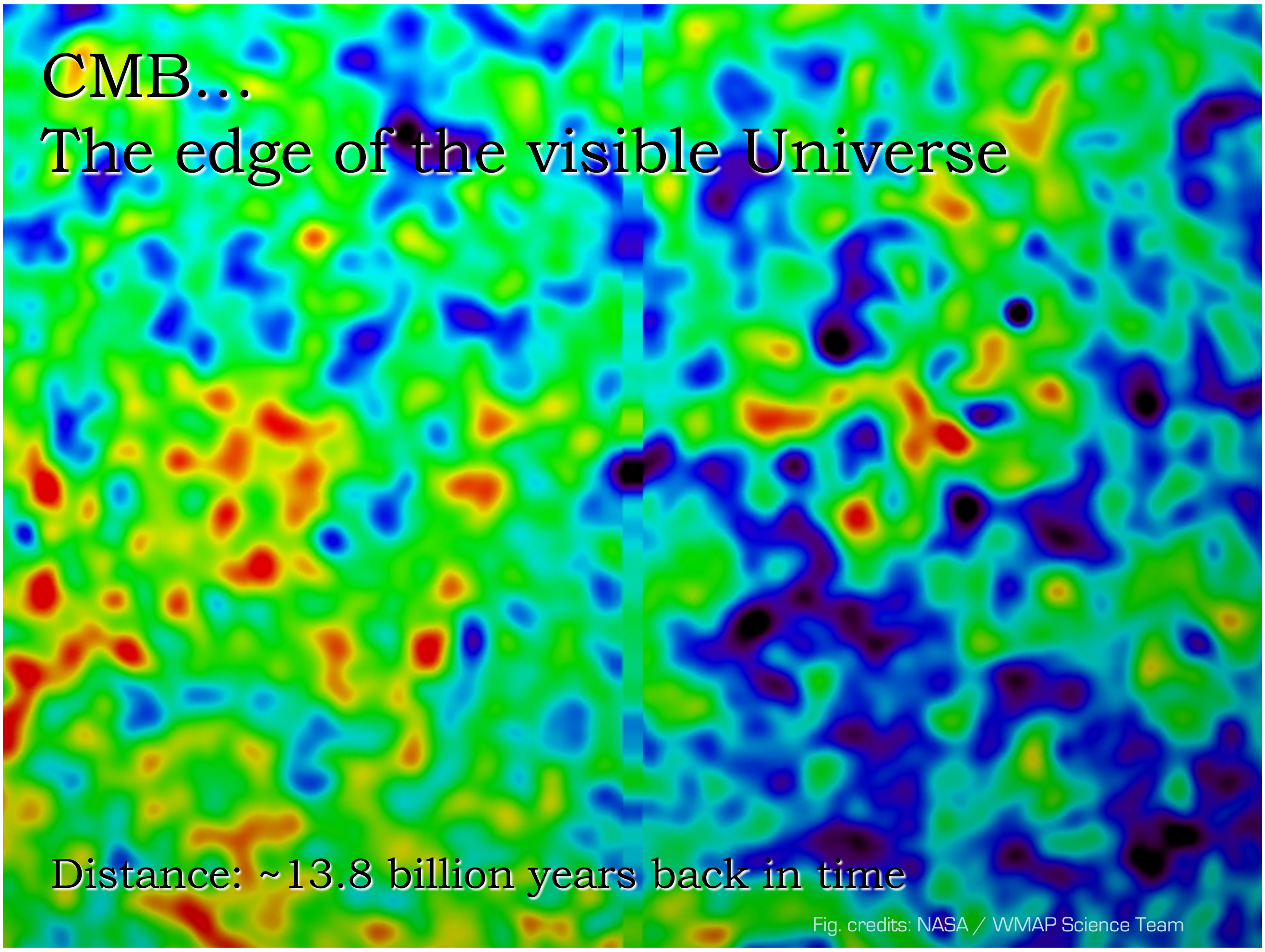
LSS: "The fingers of God"…



Credit: M. Blanton and the Sloan Digital Sky Survey

Further away and back in time…
First stars forming period





CMB...
The edge of the visible Universe

Distance: ~13.8 billion years back in time

Fig. credits: NASA / WMAP Science Team

Planck Surveyor: looking back to the dawn of time

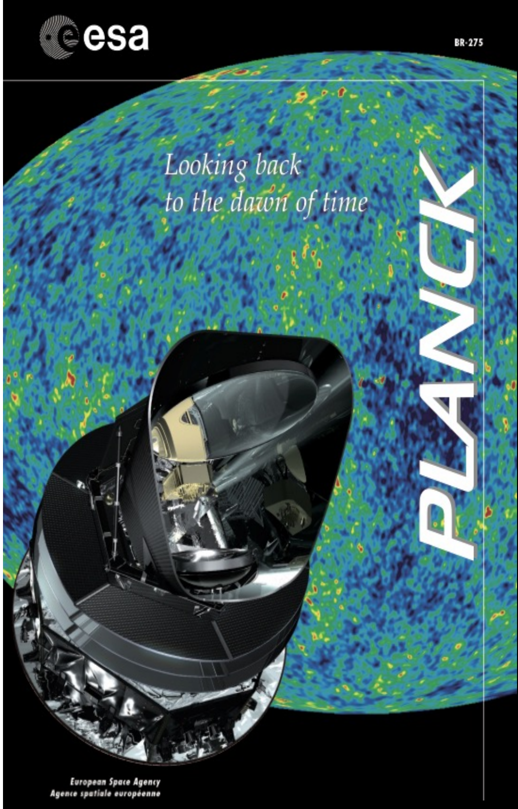


Fig. credits: ESA

Project: ESA lead mission to observe the temperature and polarization anisotropies of the Cosmic Microwave Background (CMB) radiation with unprecedented precision.

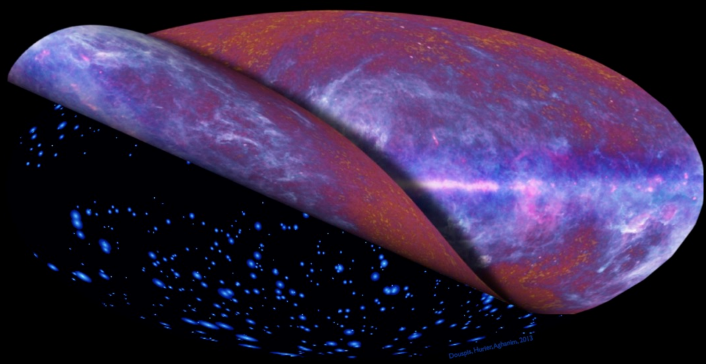
Total Cost: about €700 million (€1 / person in EU)

Mission timeline:
- Launch: 14 May 2009
- Operational orbit at L2: July 2009
- Nominal science phase: end of January 2011
- Extended mission: Shut down date: 19 Oct. 2013

Payload:
- Telescope: 1.5 m projected apertures
- Low Frequency Instrument (LFI): array of 22 tuned radio receivers operating at 30, 44 and 70 GHz.
- High Frequency Instrument (HFI): array of 52 bolometers operating at 100, 143, 217, 353, 545, and 857 GHz.

Planck CMB observations

2009-2013: Planck satellite observes the CMB sky with unprecedented angular resolution and sensitivity.



Animation credits: ESA and the Planck collaboration; Cluster map by Douspis, Hurier, Aghanim 2013

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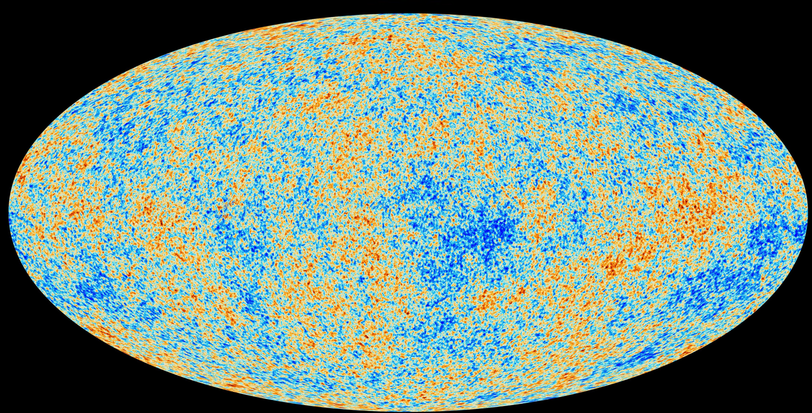
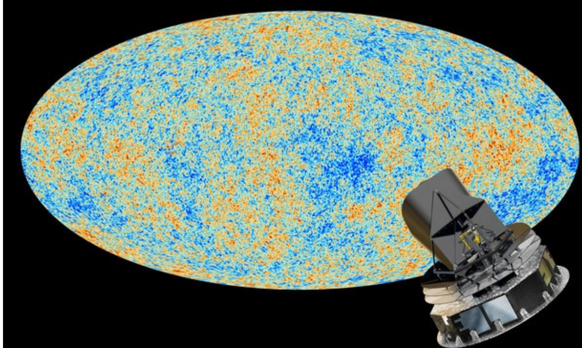


Animation & Fig. credits: ESA and the Planck collaboration

Galaxy surveys: 3D mapping of the Universe…

SDSS: aims at ~25% of the sky; ~100 million objects

ESA Cosmology missions



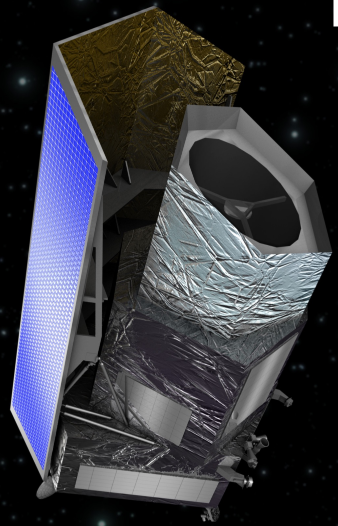
ESA/NASA Planck Surveyor:
2009 - 2013

ESA Euclid:
2023 - 2029

esa

Euclid mission (ESA): Galaxy Surveys from space (launch 2023)

Portuguese official participation lead by the cosmology group @ IA – CAUP/FCUL



Euclid: ~ 2000 million galaxies

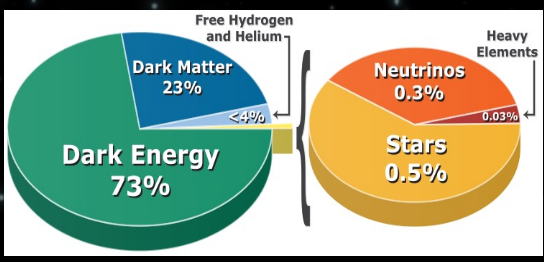
Fig. credits: ESA - C. Carreau.

Euclid mission (ESA): Galaxy Surveys from space (launch 2023)

Primary Objective: To map the distribution and evolution of (dark) matter in the Universe to:

- Measure the accelerated expansion
- Unveil the nature of dark energy
- Test models of gravity, dark matter, and of the primordial universe

Present observations suggest that dark energy (DE) and dark matter (DM) are 96% of the total energy density budget of the universe.



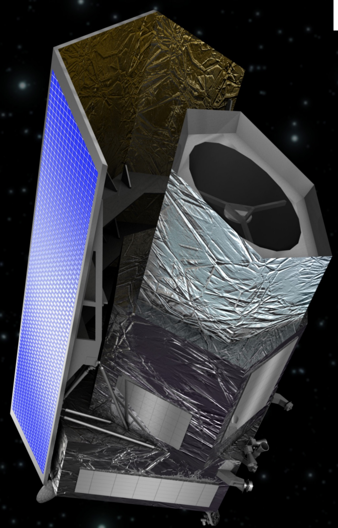
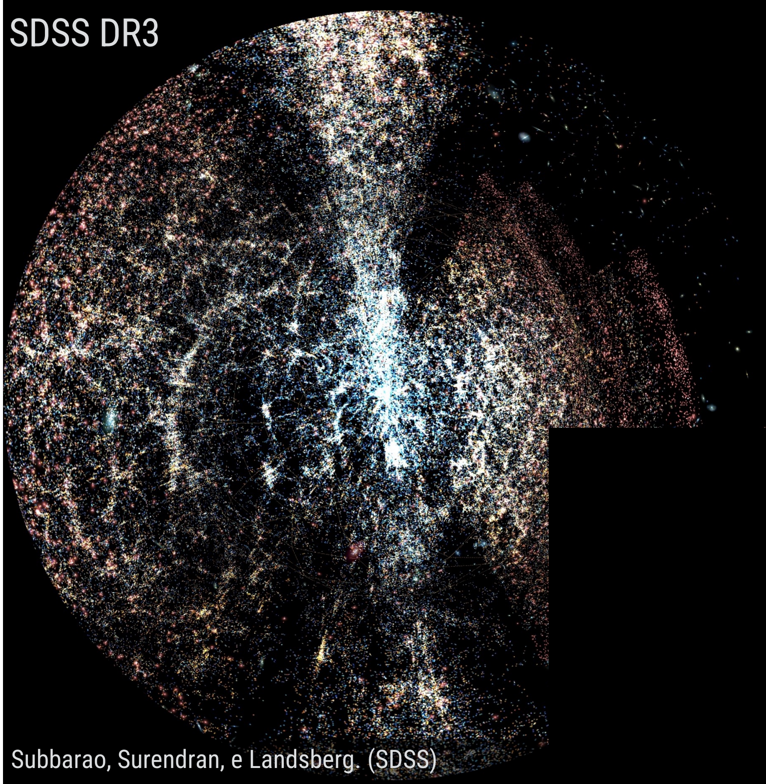
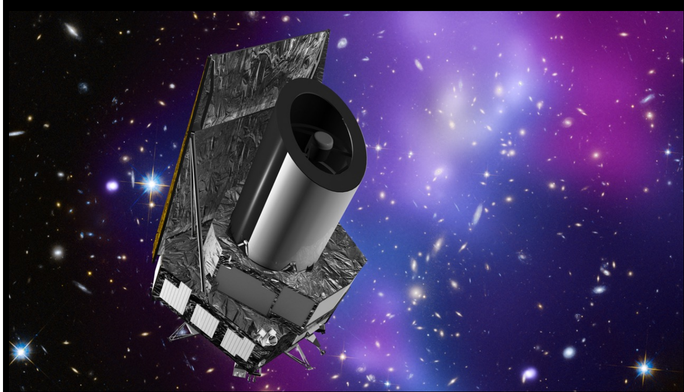


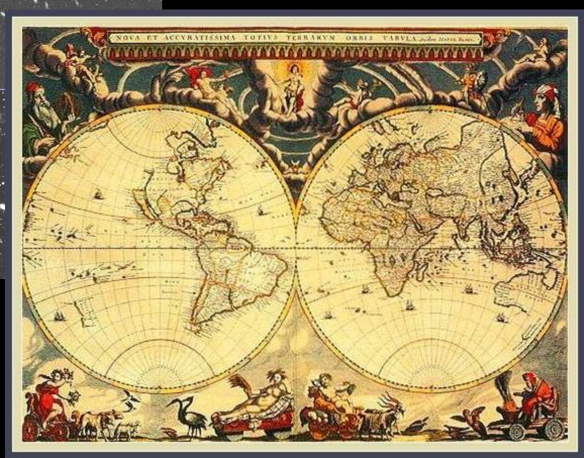
Fig. credits: ESA - C. Carreau.

Levantamento de galáxias em 3D

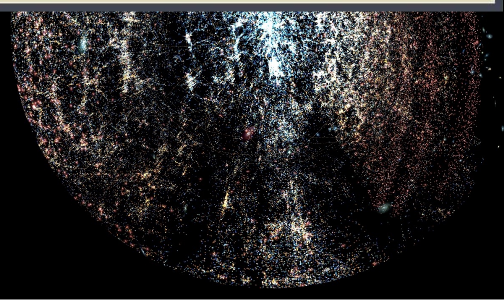


Cartografia de galáxias em 3D com o Euclid

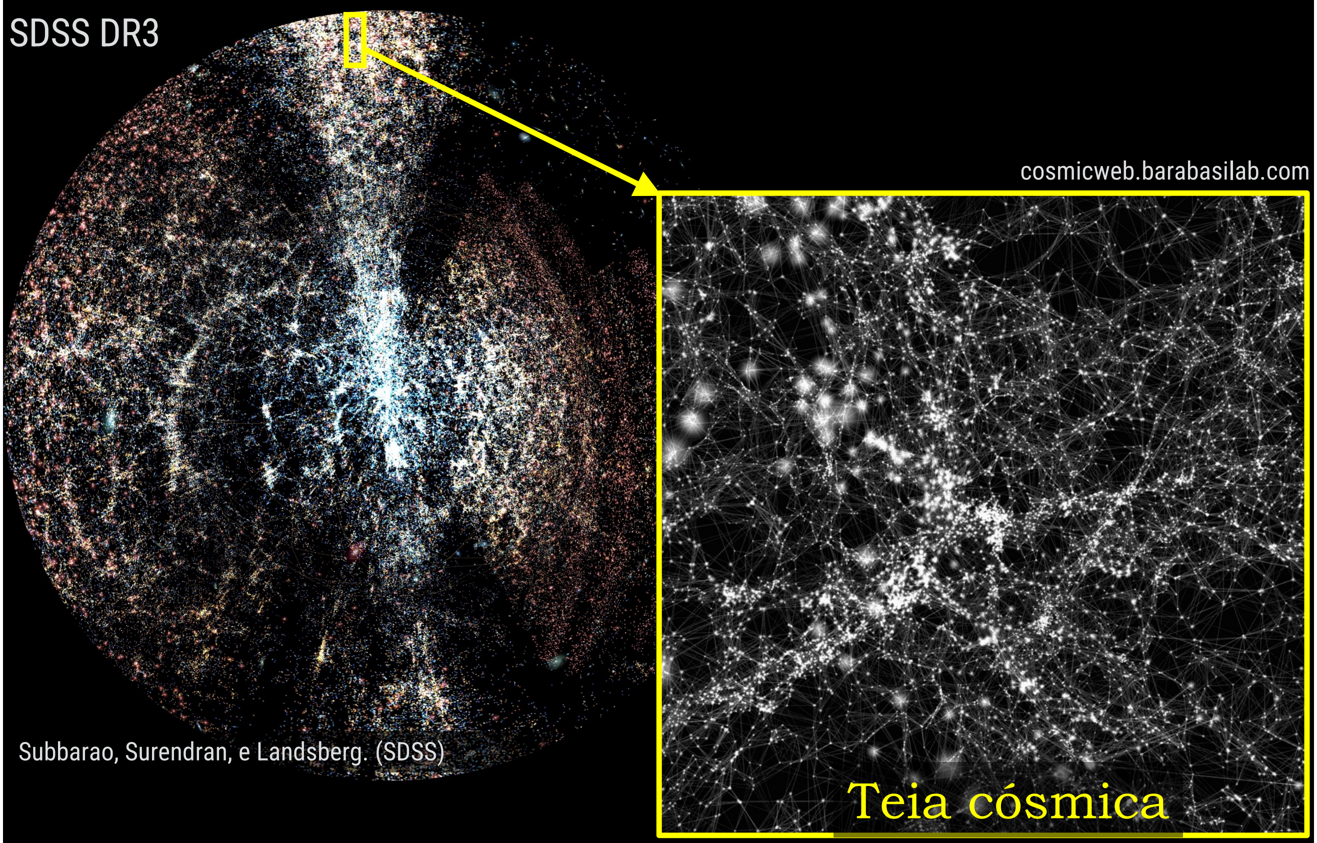




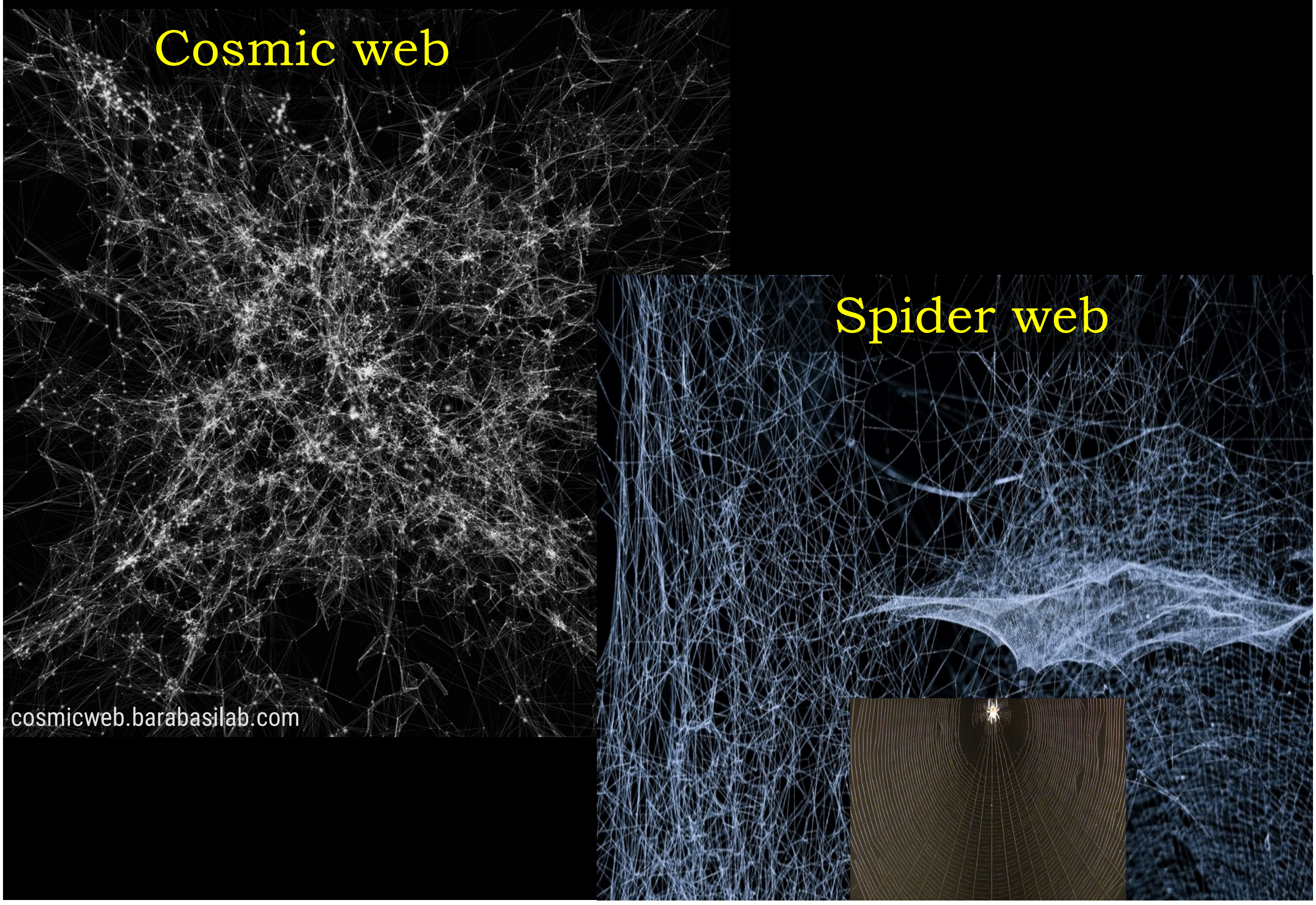
O Euclid fará a maior e mais precisa cartografia, de galáxias no Universo, observando imagens e obtendo espectros a partir do espaço.



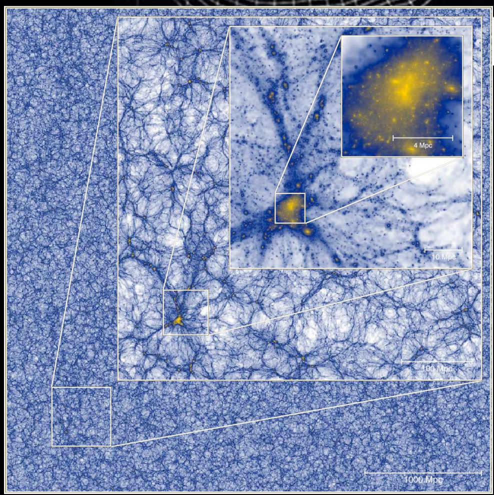
Como se determina a natureza da energia escura?



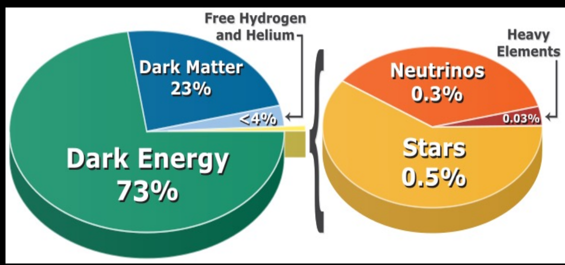
Anology with the study of living organisms



Analogia com o estudo dos seres vivos

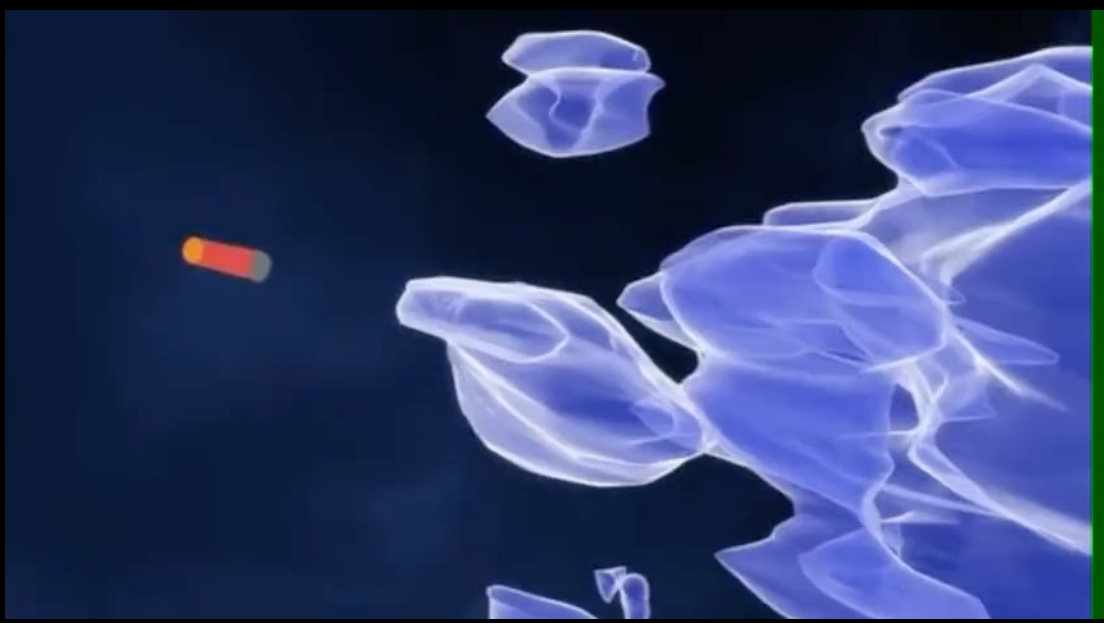


A teia cósmica é muito sensível a parâmetros Cosmológicos e de Física fundamental



Euclid: missão classe M do programa científico da ESA

O Euclid vai observar mais de mil milhões de galáxias durante 6 anos. A luz emitida pelas galáxias é deflectida por estruturas invisíveis de matéria escura através do efeito de lentes gravitacionais, sendo observadas com distorção.



Euclid: missão classe M do programa científico da ESA

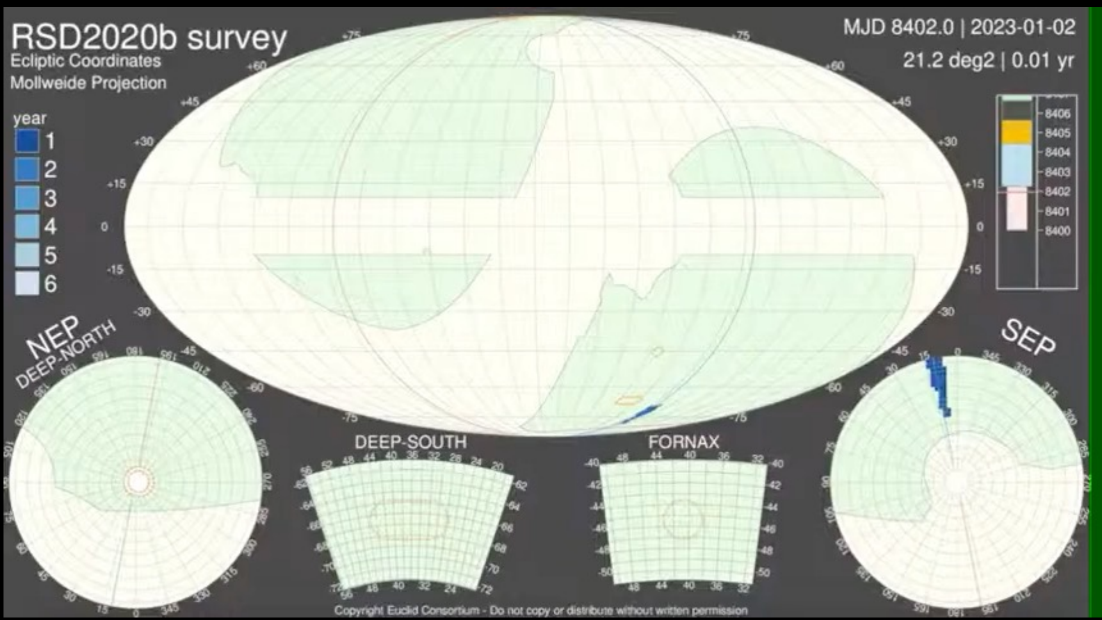
As distorções são muito pequenas. Mas com a observação de mil milhões de galáxias há estatística suficiente para mapear a matéria escura e estudar a natureza da Energia Escura.

Imagem do Hubble do enxame de galáxias Abell 370
(a 4 biliões de anos luz da terra)

Euclid: missão classe M do programa científico da ESA

Para se obter boas imagens o satélite vai ser colocado em orbita do ponto de Lagrange L2, de "costas voltadas" para o sol, apontando sucessivamente para diferentes direções das regiões mais escuras do céu. Cada campo terá de ser observado durante 1h15. A área total de observação será de 15000 graus quadrados.

Euclid Sky Survey implementation and optimization is a PT, national, responsibility

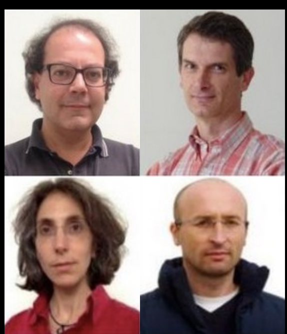


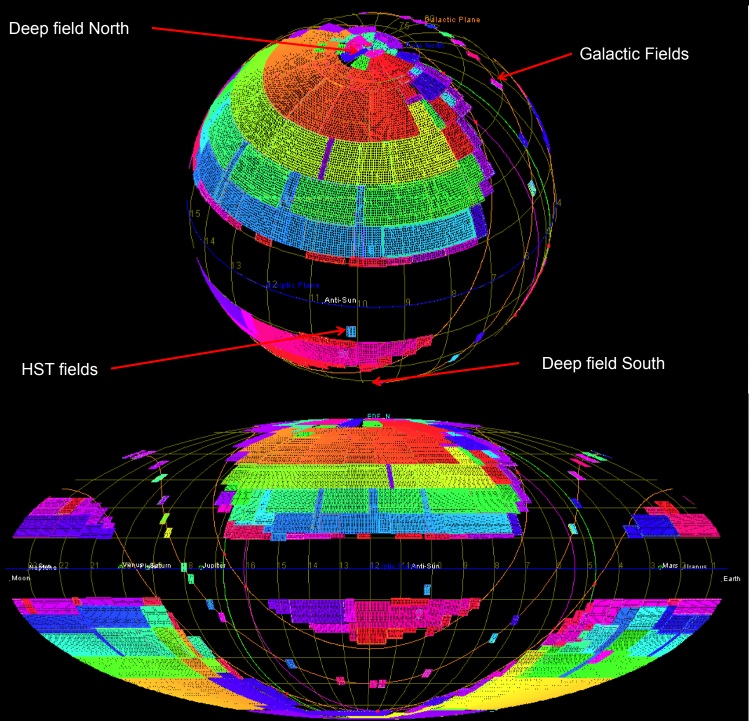
Animação criada por J. Dinis (FCUL / IA)

Euclid Survey planning

Portuguese official contribution is carried out at IA/FCUL:

- I. Tereno
- J. Dinis
- C. S. Cavalho
- A. da Silva





From: Euclid Preparation I. The Euclid Reference Survey: status at the Preliminary Design Review, submitted to the ECEB, 2019

Observational cosmology: empirical facts about the Universe

1. The Universe is expanding

Edwin Hubble

1912: Vesto Slipher is the first to observe spectral line (red)shifts towards nebulae and to relate these redshifts to their recessional velocities.

1924: Edwin Hubble ends debate on the nature of nebulae being galactic objects

1929: reports a linear relation between relative radial velocity and distance: $v = Hd$

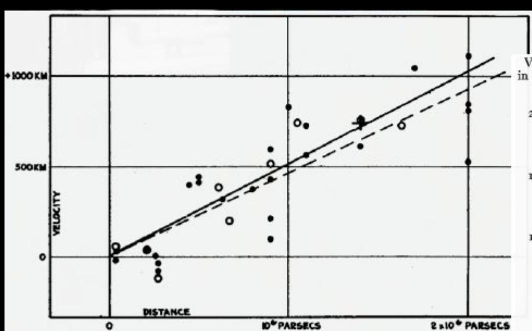
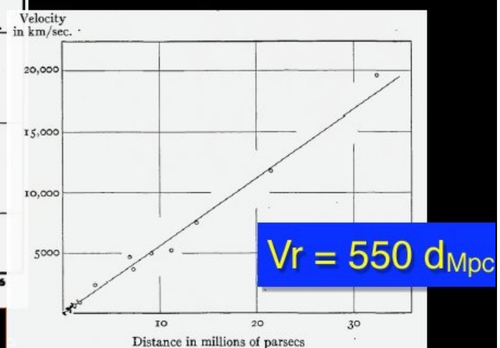


FIGURE 1



From:

1. The Universe is expanding

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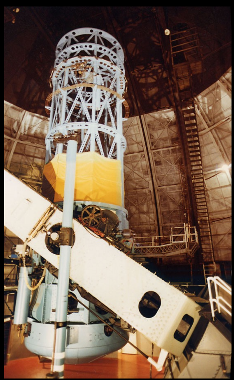

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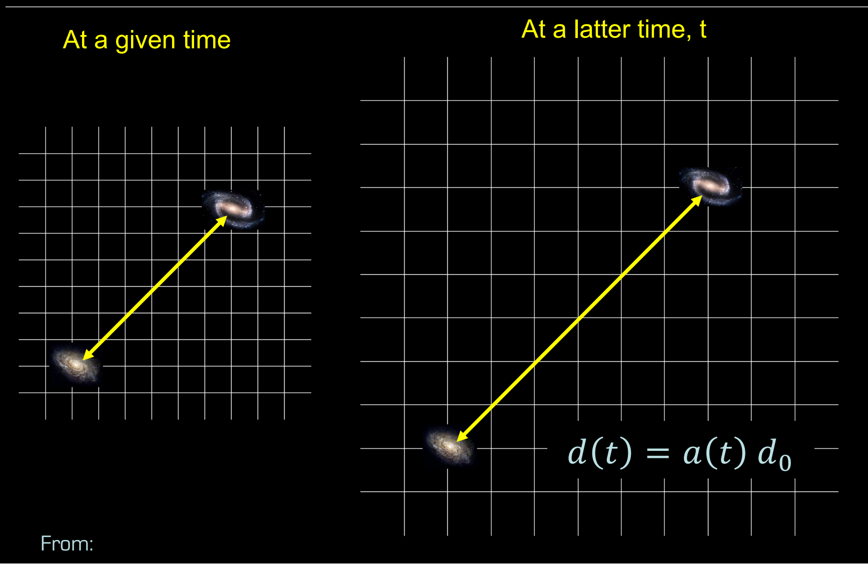

From:

1. The Universe is expanding

Edwin Hubble

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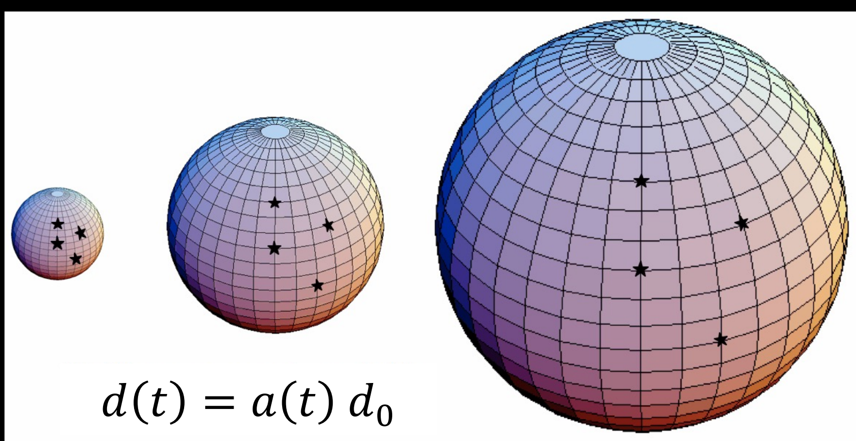
Time evolution of an expanding spherical surface



From:

The basic idea behind the Big-Bang theory

- If the universe is expanding and matter-energy is conserved during the expansion then the universe had to be smaller, denser and hotter in the past!

- If so, the Universe must have evolved from a state where matter and radiation form a ultra dense and hot ionized plasma of fundamental particles

- As the universe expands and cools down:
    - interactions between the plasma components become less frequent;
    - different particle species should decouple from the plasma;
    - eventually the universe becomes neutral and transparent to radiation

According to the Big-Bang theory, in the early instants…

“ the Universe was a extremely hot and dense plasma, like a ‘torrid bright fog’…

… radiation was trapped in this plasma through collisions with other plasma particles

… as the universe expands, the plasma temperature drops, atomic nuclei form, and capture the free electrons in the plasma. When the number of free electrons is too small, radiation no longer interacts with the plasma and propagates freely, giving rise to the Cosmic Microwave Background and neutral matter”

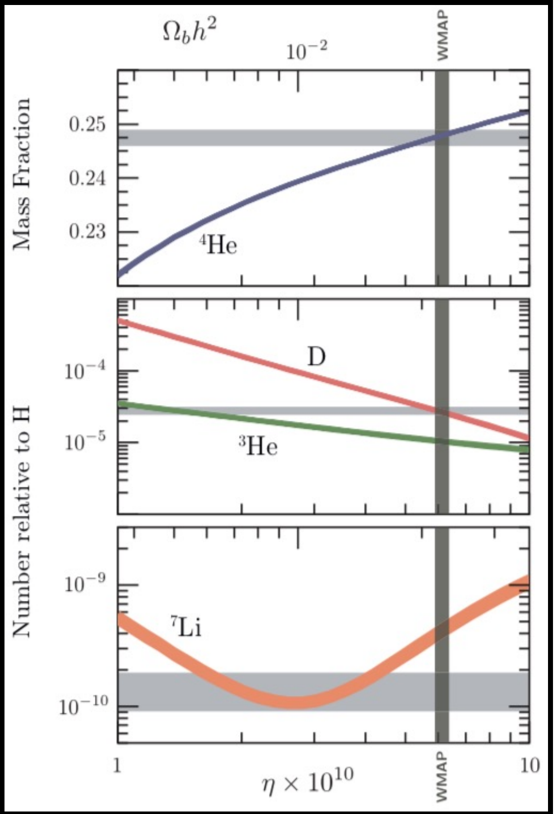
2. The abundance of light nuclei

Herman, Gamow, Alpher

The relative abundance of light elements can not be explained by stellar nucleosynthesis

1948: Alpher & Gamow computed the abundance of light elements in the context of the Big Bang theory

Light elements were produced at low temperatures (<1e9K and high densities) during several tens of minutes



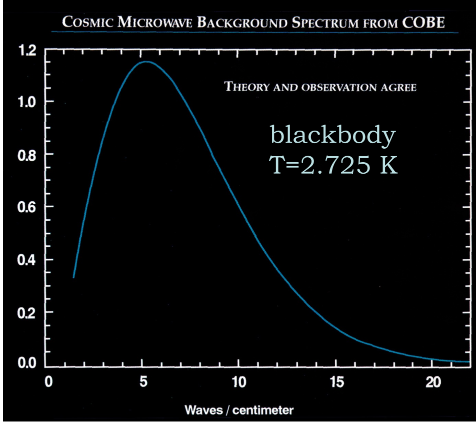
From:

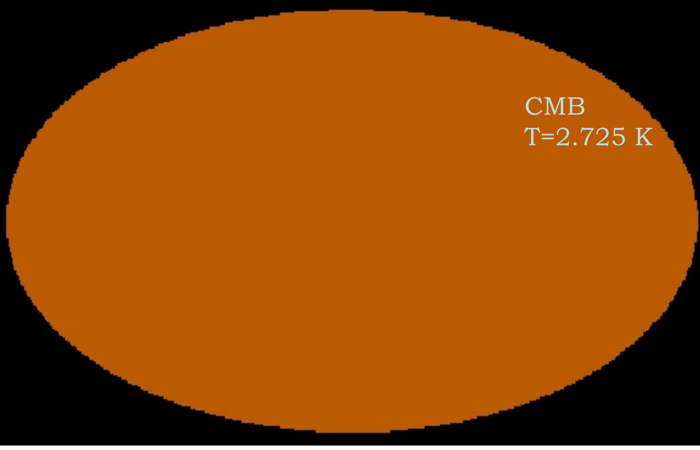
3. Cosmic Microwave Background

Penzias & Wilson

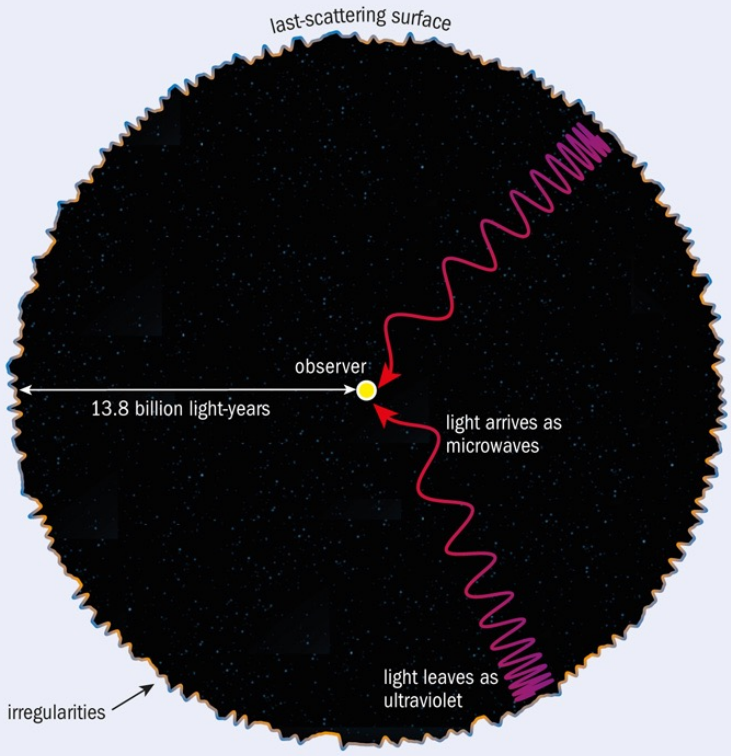
1965: Penzias & Wilson serendipitously discovered a uniform radiation ("excess") across the sky.

This was the cosmic microwave background radiation predicted by Gamow and Alpher in 1948





CMB: the last scattering surface



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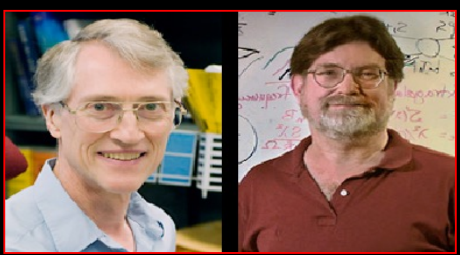
Reprinted from: http://physicsworld.com/cws/article/indepth/2014/jan/09/planck-perspectives

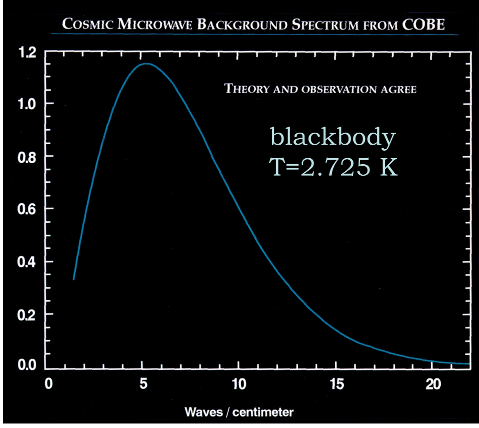
3. Cosmic Microwave Background



John Matter & George Smooth

1991: High precision measurement of CMB temperature by COBE and 1st detection of temperature fluctuations (Mather & Smoot)

2001: State of the art measurements of dT/T~1e-5 temperature fluctuations by WMAP



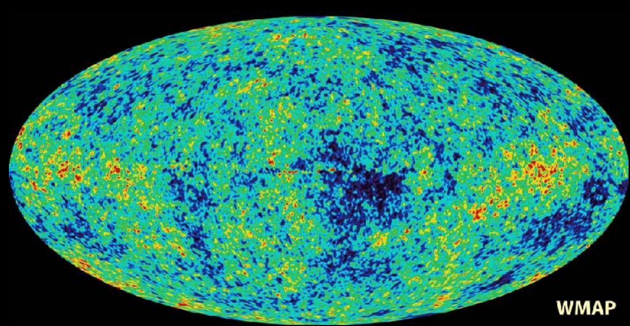
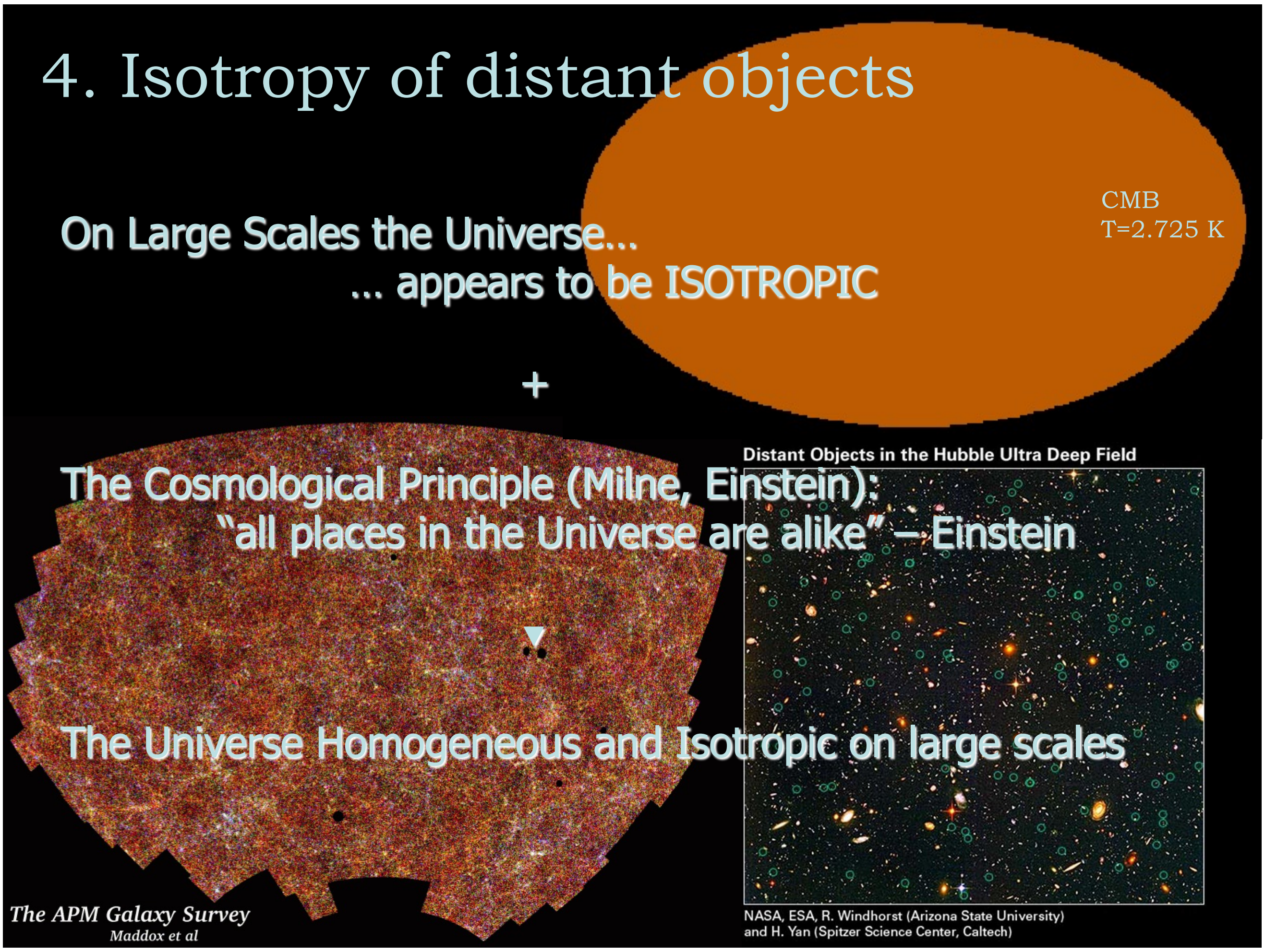


Fig. credits: NASA / WMAP Science Team



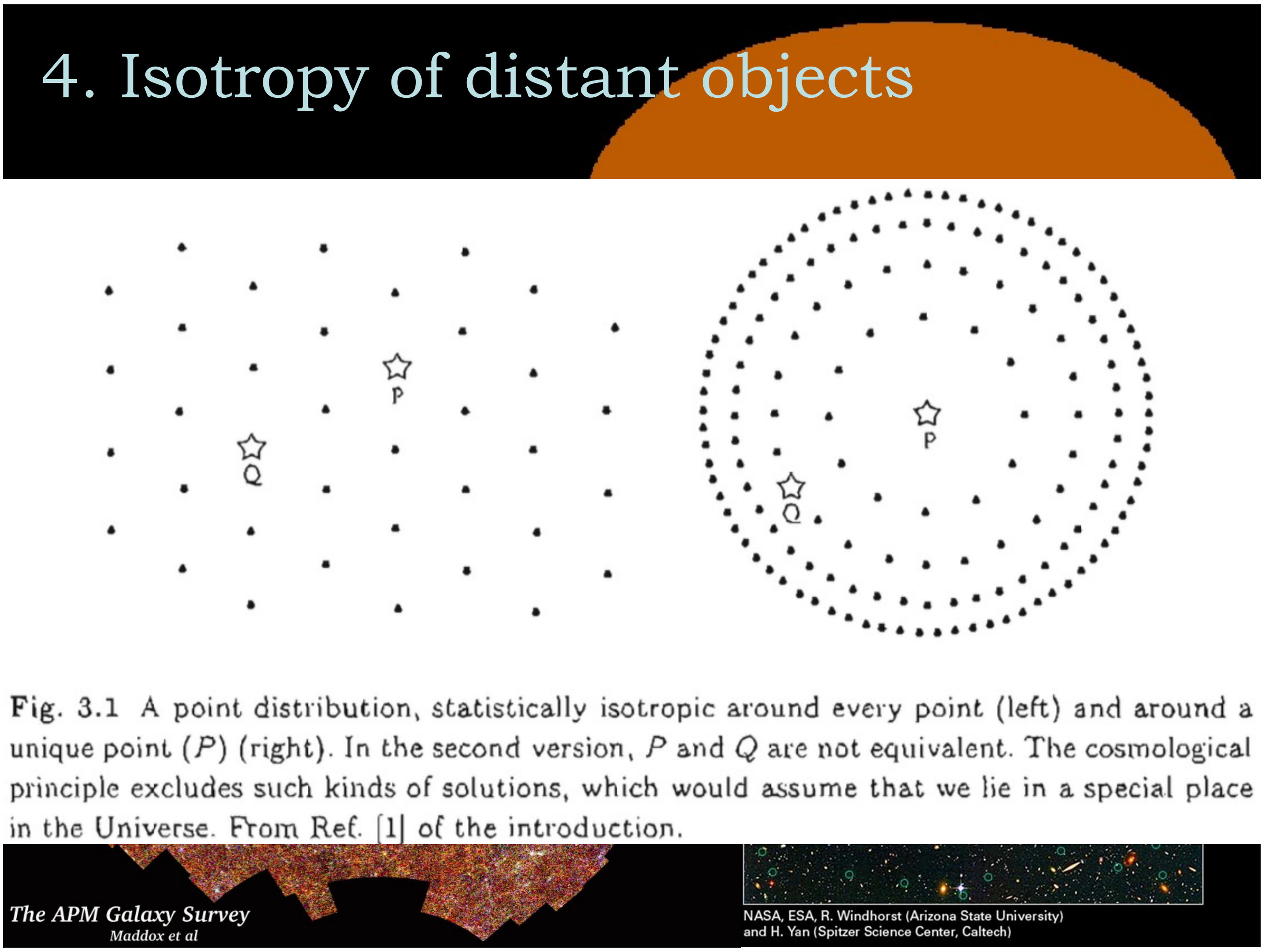
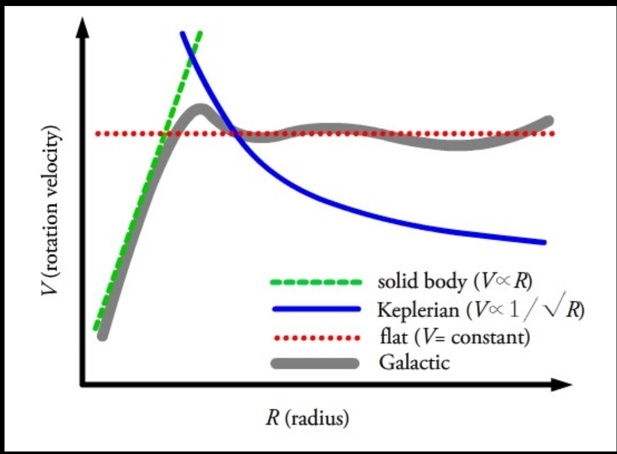


Fig. 3.1 A point distribution, statistically isotropic around every point (left) and around a unique point (P) (right). In the second version, P and Q are not equivalent. The cosmological principle excludes such kinds of solutions, which would assume that we lie in a special place in the Universe. From Ref. [1] of the introduction.

5. The existence of Dark Matter



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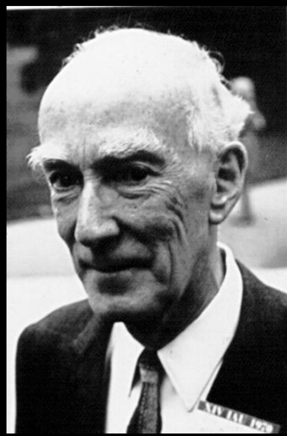


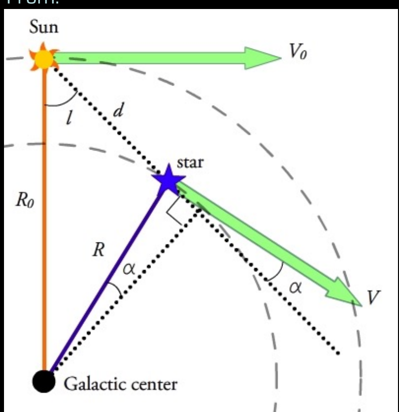
Jan Oort

1927: Jan Oort studies the rotation of stars in our galaxy and infers that their rotation is not consistent with Keplerian motion.

$$v_{circ} = \sqrt{\frac{GM(R)}{r}}$$

5. The existence of Dark Matter

From:



Circular motion:

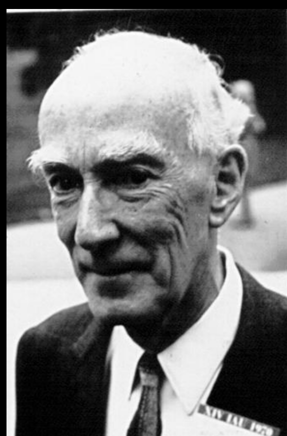
$$v_{circ} = \sqrt{\frac{GM(R)}{r}}$$

If the whole mass is mostly at the centre: v_{cir} ^ 2 ~ 1/r

Oorts constants:

$$A \equiv -\frac{1}{2}\left[\frac{dV_c}{dR}\Big|_{R_0} - \frac{V_{c,0}}{R_0}\right]$$

$$B \equiv -\frac{1}{2}\left[\frac{dV_c}{dR}\Big|_{R_0} + \frac{V_{c,0}}{R_0}\right]$$



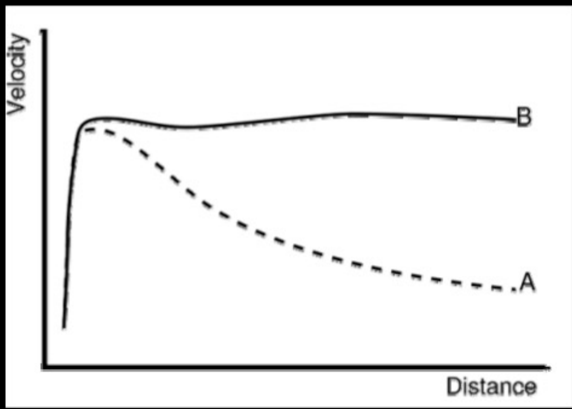
Observations vs Keplerian motion:

- Kepler. motion: (A-B)/(A+B) = 2
- Observations : (A-B)/(A+B) = 5

-Mass is not concentrated at the centre
-Non-luminous mass is required

http://icc.dur.ac.uk/~tt/Lectures/Galaxies/TeX/lec/node42.html

5. The existence of Dark Matter



B: Observations

A: theoretical expectations

1980: Vera Rubin and others also find that stars rotate too fast in the outskirts of spiral galaxies to remain bound assuming that gravity is produced only by visible matter.

5. The existence of Dark Matter

Fritz Zwicky

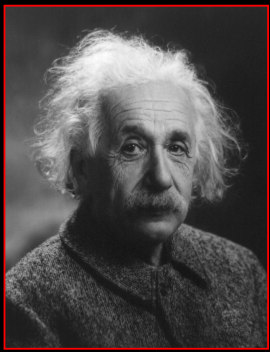
1936: Fritz Zwicky applied the Virial theorem to the velocities of galaxies in the Coma cluster and finds very high mass-to-light ratios, $\Upsilon = M/L$, for them to remain bound: $\Upsilon_{coma}/\Upsilon_{sun} = 500 \gg$ 2-10 for galaxies.

- Virial theorem (for gravitationally relaxed systems): $2\bar{E}_k + \bar{E}_p = 0$
- Mass from the virial theorem: $M_V =< v^2 >< R >/G$
- Visible luminous Mass: $M_L = N_g\, \Upsilon_g\, L_g$

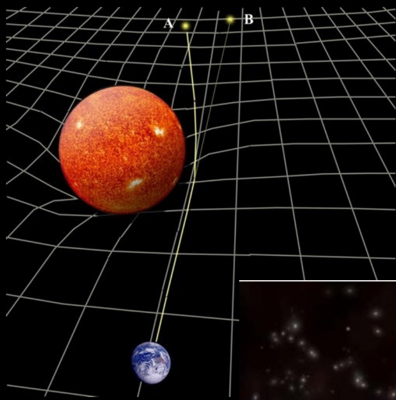
(N_g - number of galaxies; Υ_g - galaxy mass-to-light ratio; L_g galaxy luminosity)

5. The existence of Dark Matter

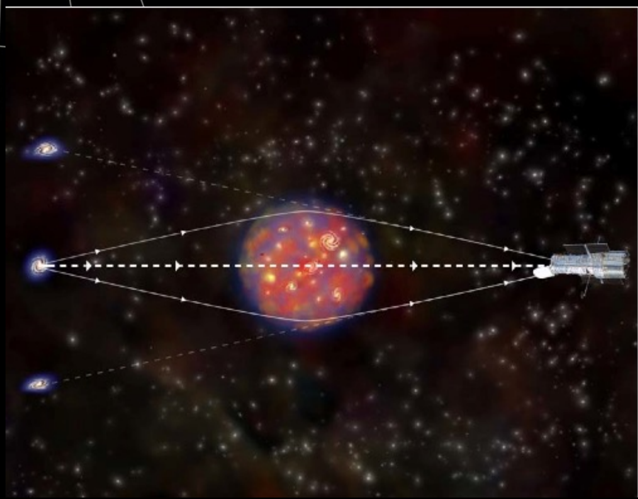
lensing effects:



Albert Einstein



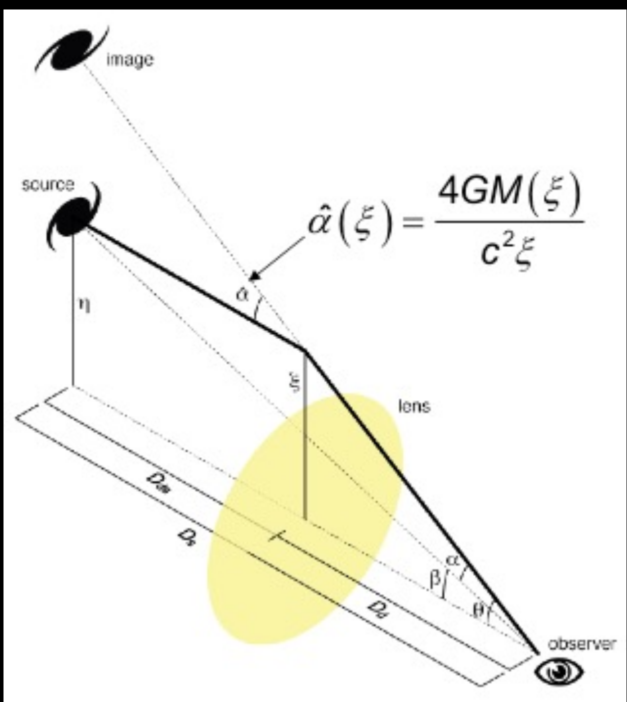
$$G_{\mu\nu} = 8\pi G T_{\mu\nu}$$



From:

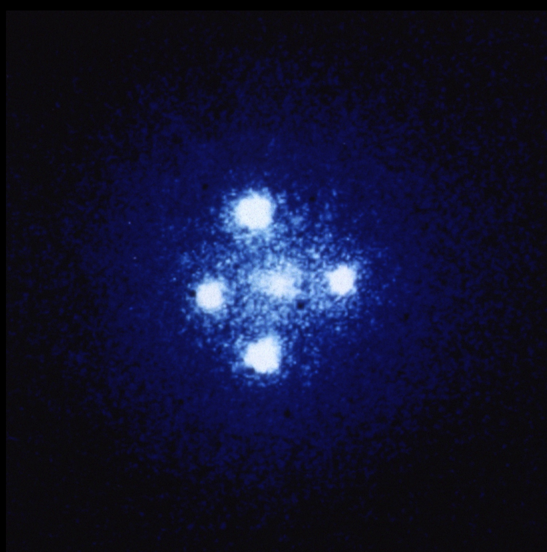
5. The existence of Dark Matter

lensing effects:



From:

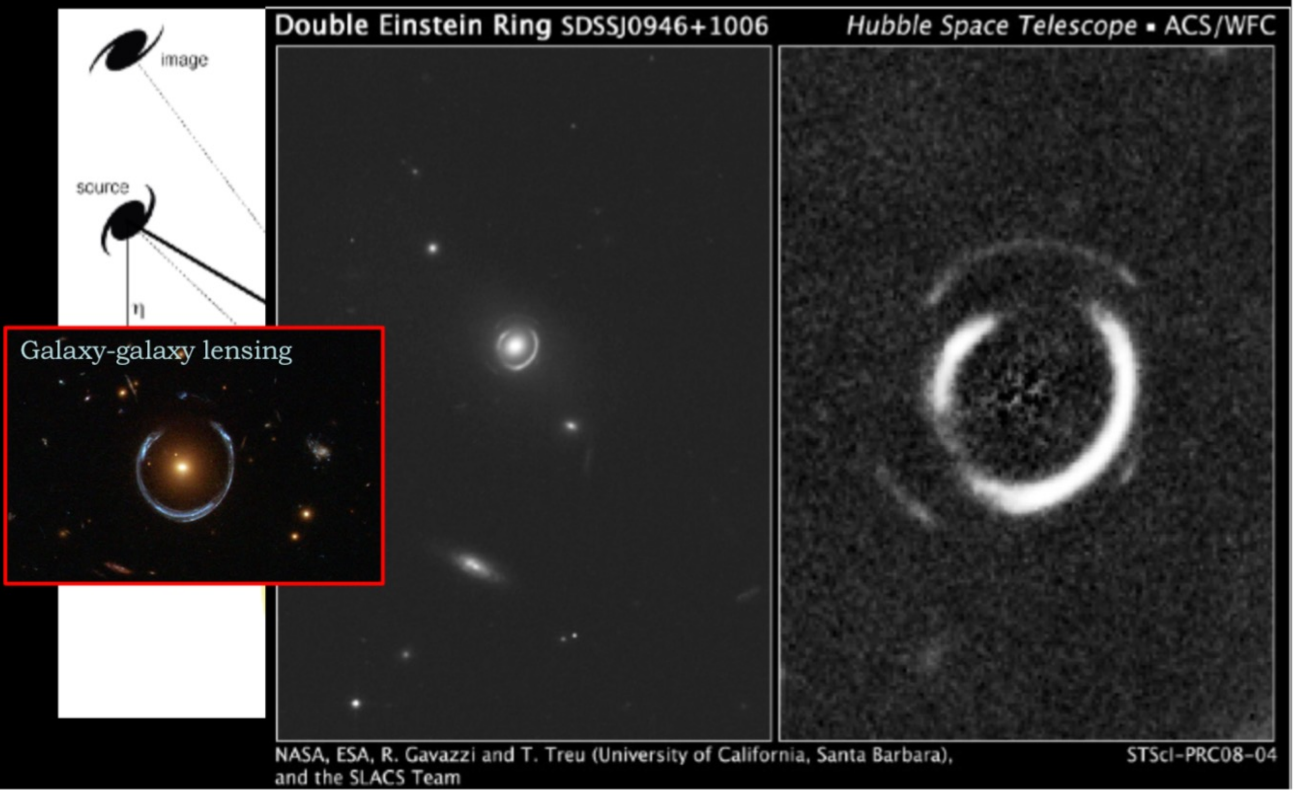
Strong lensing



5. The existence of Dark Matter

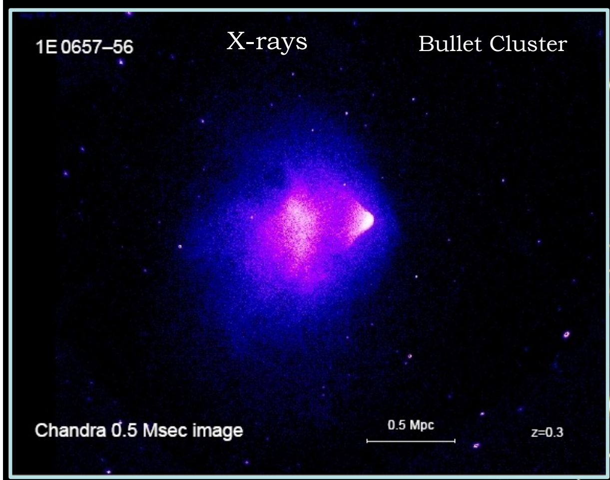
lensing effects: strong lensing

Einstein Rings

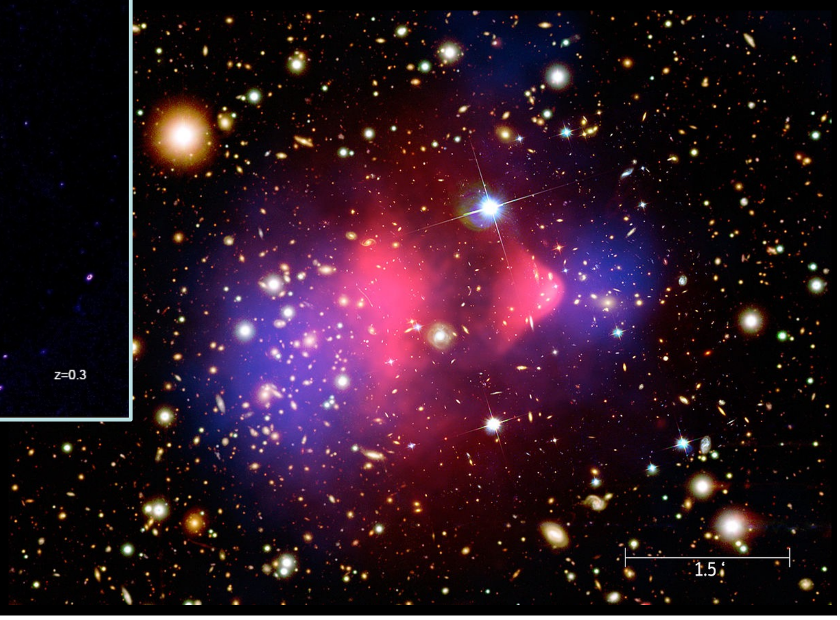


5. The existence of Dark Mater

2003: X-ray (produced by extremely hot gas – in red) vs weak lensing observations (probing the total mass distribution in blue) of the Bullet Cluster put in evidence that galaxy clusters must contain "dark matter"



X-rays + lensing + visible



Refs:
astro-ph/0309303
astro-ph/0312273

6. Cosmic expansion is accelerating

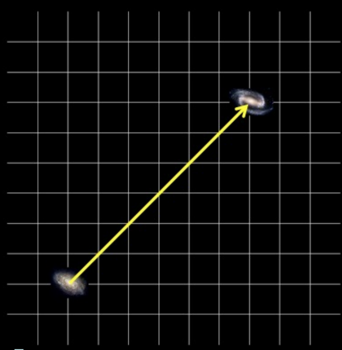
1998: S. Perlmutter and the supernova Cosmology project found first evidence for the accelerated expansion of the Universe.

assuming supernovae are standard candles, they appear further away (green arrow) then predicted by non-accelerating expansion models (yellow arrow).

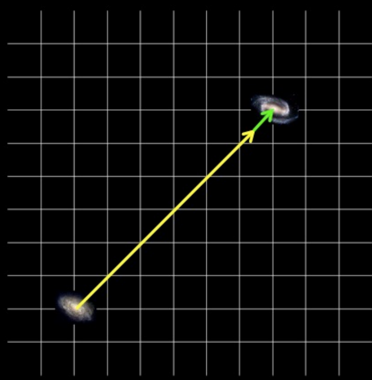
$$d(t) = a(t)\, d_0$$

with $\ddot{a}(t) > 0$

No Accelerated expansion

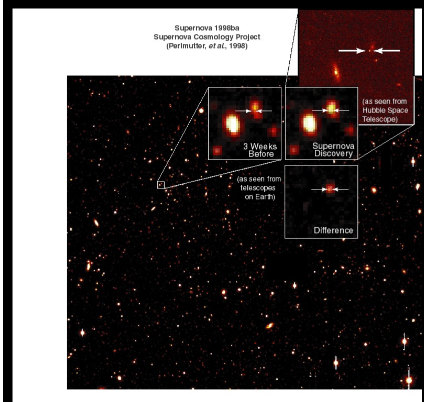


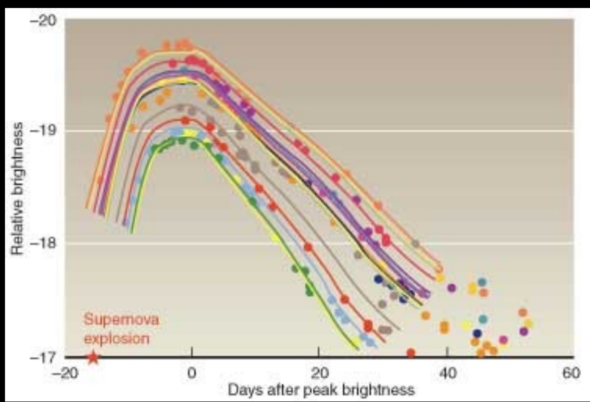
Accelerated expansion



From:

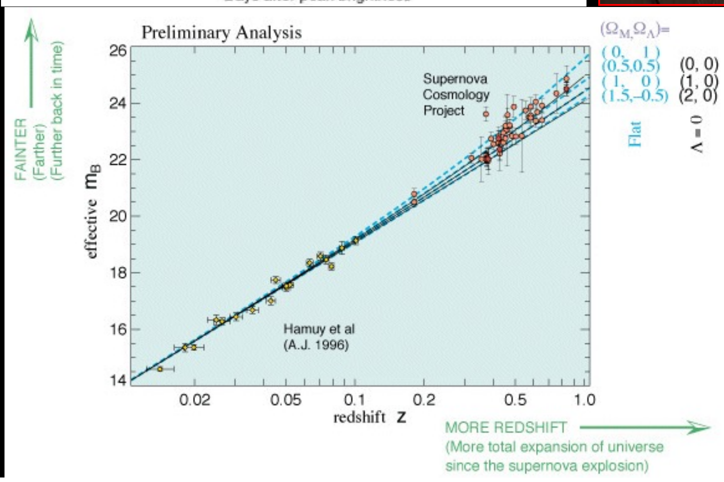
6. Cosmic expansion is accelerating



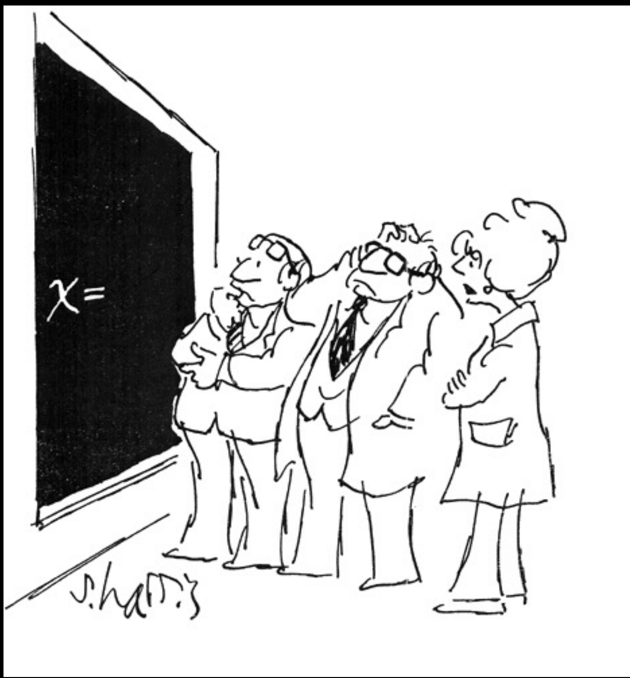


Cosmological redshift:

$$z = \frac{E - E_0}{E_0} = $$
$$= \frac{\nu}{\nu_0} - 1 =$$
$$= \frac{\lambda_0}{\lambda} - 1 =$$
$$= \frac{a_0}{a} - 1$$



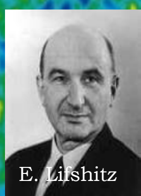
How Cosmological structure forms and evolves?



- ❑ Observations indicate that
  - ❑ on small scales the universe is NOT homogeneous and isotropic
  - ❑ On large cosmological scales the Universe does not show indications of strong anisotropies. Together with the cosmological principle this implies the universe is highly homogeneous and isotropic
  - ❑ However it shows small anisotropies in the CMB.

How structure forms and evolves?

Structures should grow due to Gravitational Instability (Jeans, Lifshitz,…)

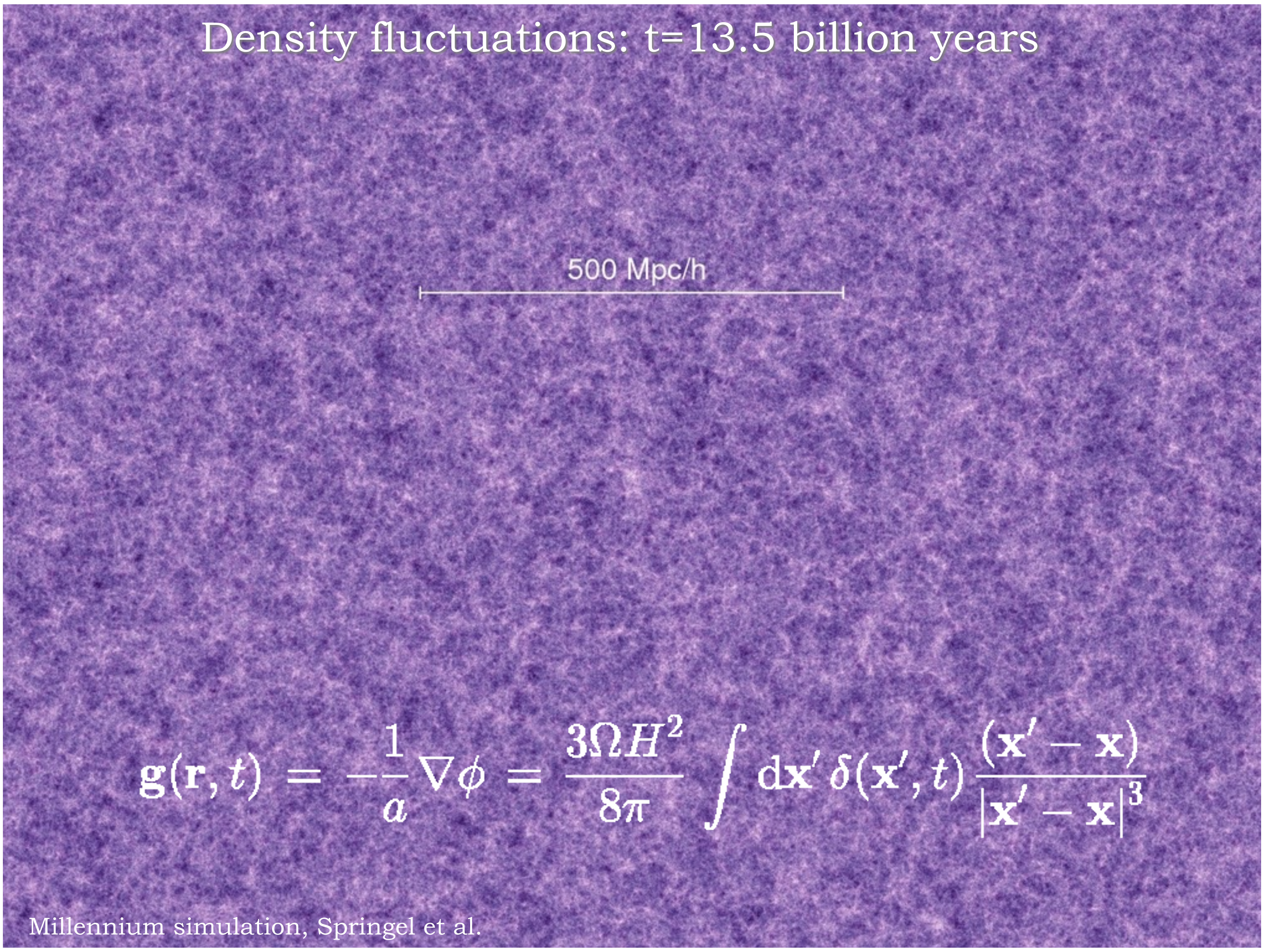


$$\mathbf{g}(\mathbf{r},t) = -\frac{1}{a}\nabla\phi = \frac{3\Omega H^2}{8\pi}\int \mathrm{d}\mathbf{x}'\,\delta(\mathbf{x}',t)\,\frac{(\mathbf{x}'-\mathbf{x})}{|\mathbf{x}'-\mathbf{x}|^3}$$

Density fluctuations: t=13.7 billion years

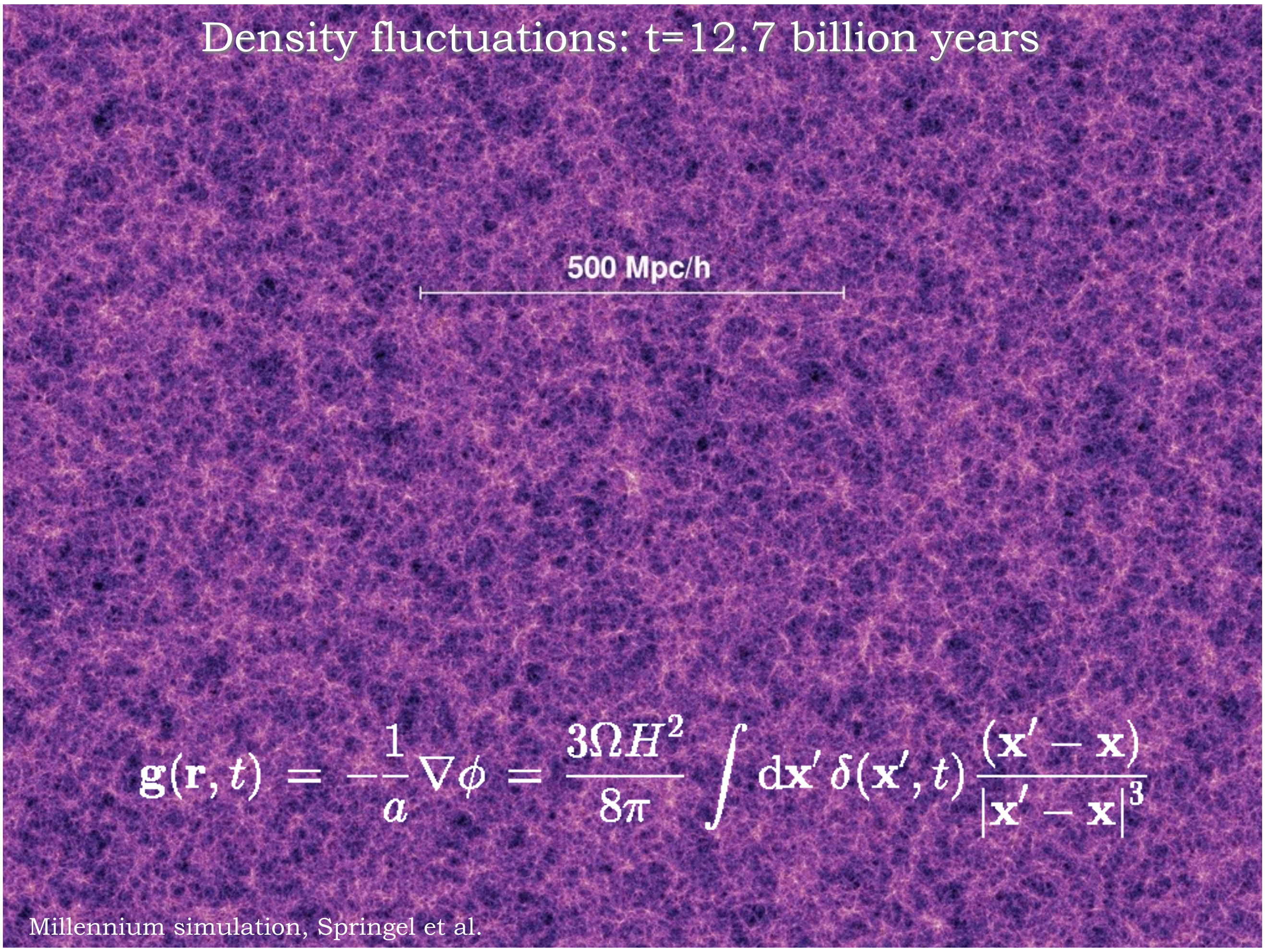
$$\mathbf{g}(\mathbf{r},t) = -\frac{1}{a}\nabla\phi = \frac{3\Omega H^2}{8\pi}\int \mathrm{d}\mathbf{x}'\,\delta(\mathbf{x}',t)\,\frac{(\mathbf{x}'-\mathbf{x})}{|\mathbf{x}'-\mathbf{x}|^3}$$

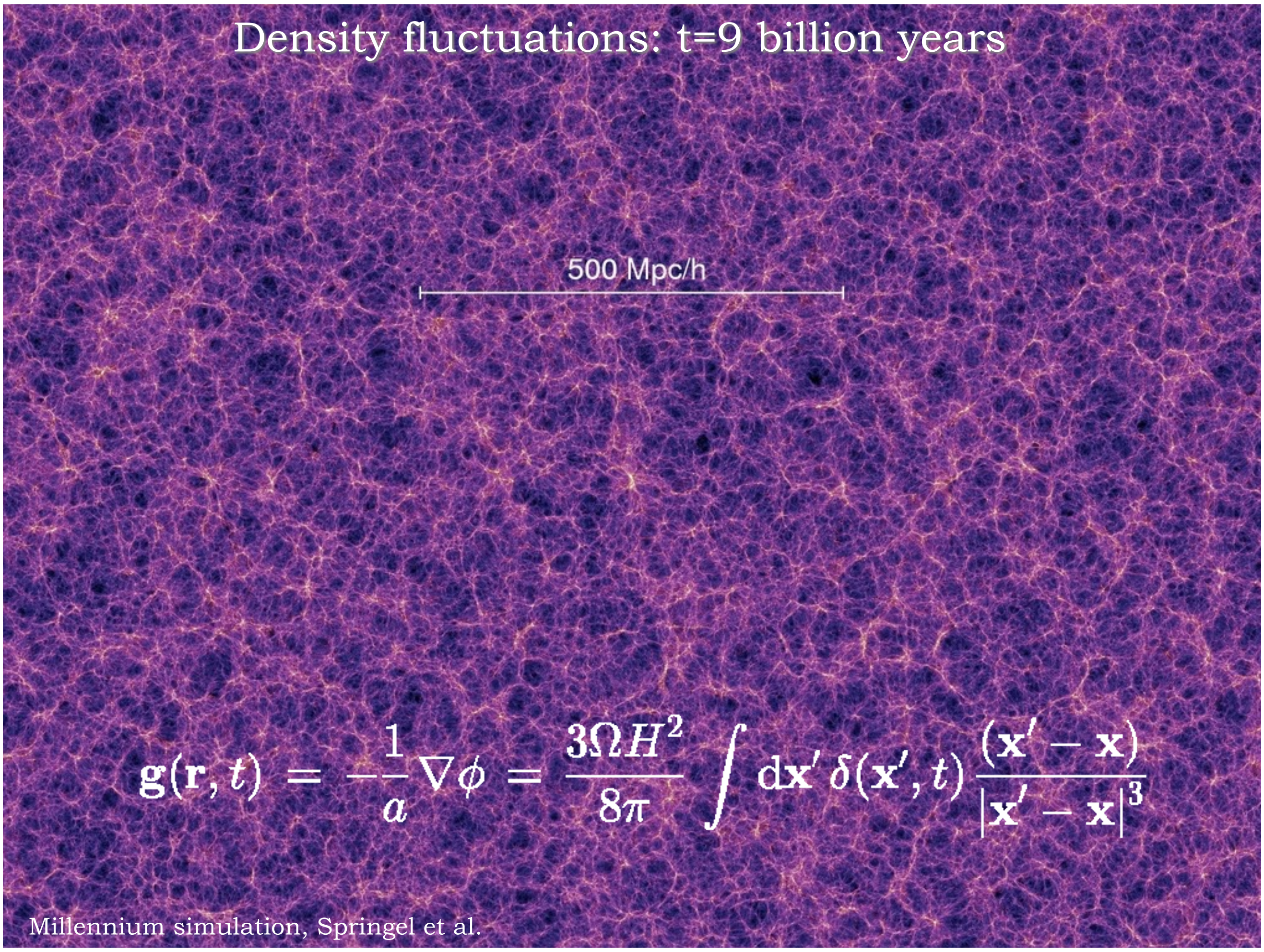
Density fluctuations: t=13.5 billion years

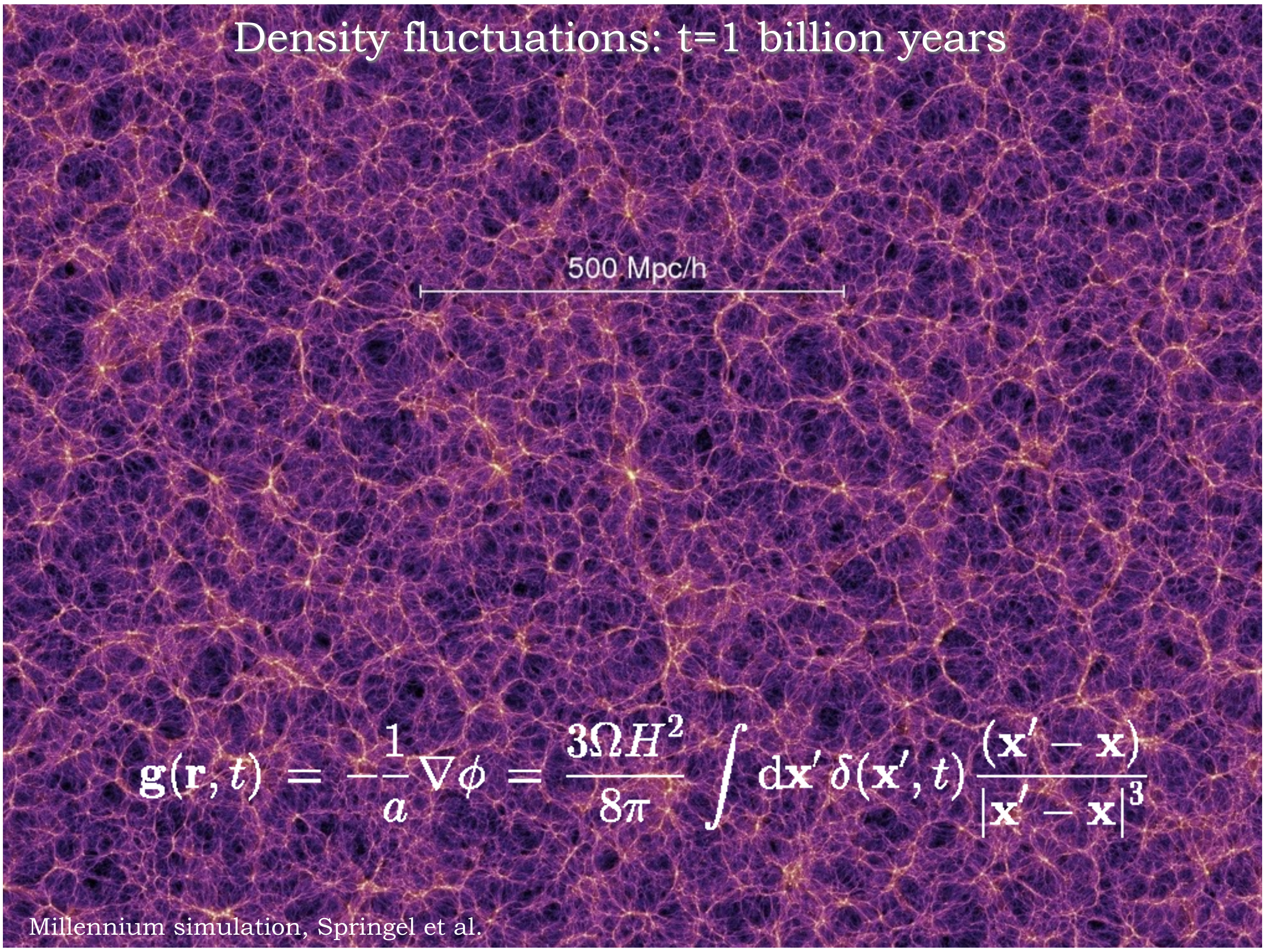


$$\mathbf{g}(\mathbf{r},t) = -\frac{1}{a}\nabla\phi = \frac{3\Omega H^2}{8\pi}\int \mathrm{d}\mathbf{x}'\,\delta(\mathbf{x}',t)\,\frac{(\mathbf{x}'-\mathbf{x})}{|\mathbf{x}'-\mathbf{x}|^3}$$

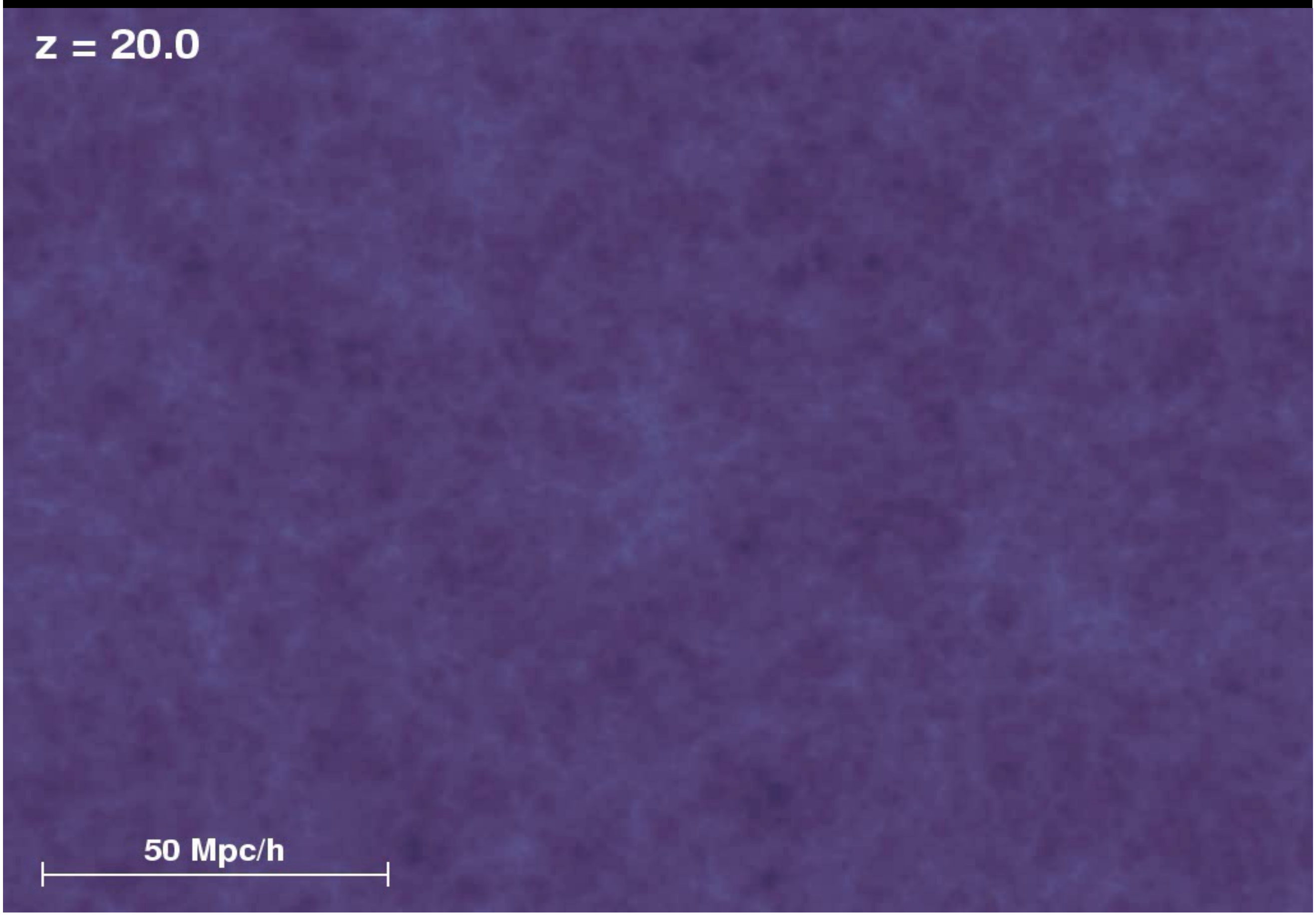
Millennium simulation, Springel et al.


Density fluctuations: t=12.7 billion years
500 Mpc/h
$$\mathbf{g}(\mathbf{r},t) = -\frac{1}{a}\nabla\phi = \frac{3\Omega H^2}{8\pi}\int \mathrm{d}\mathbf{x}'\,\delta(\mathbf{x}',t)\,\frac{(\mathbf{x}'-\mathbf{x})}{|\mathbf{x}'-\mathbf{x}|^3}$$
Millennium simulation, Springel et al.

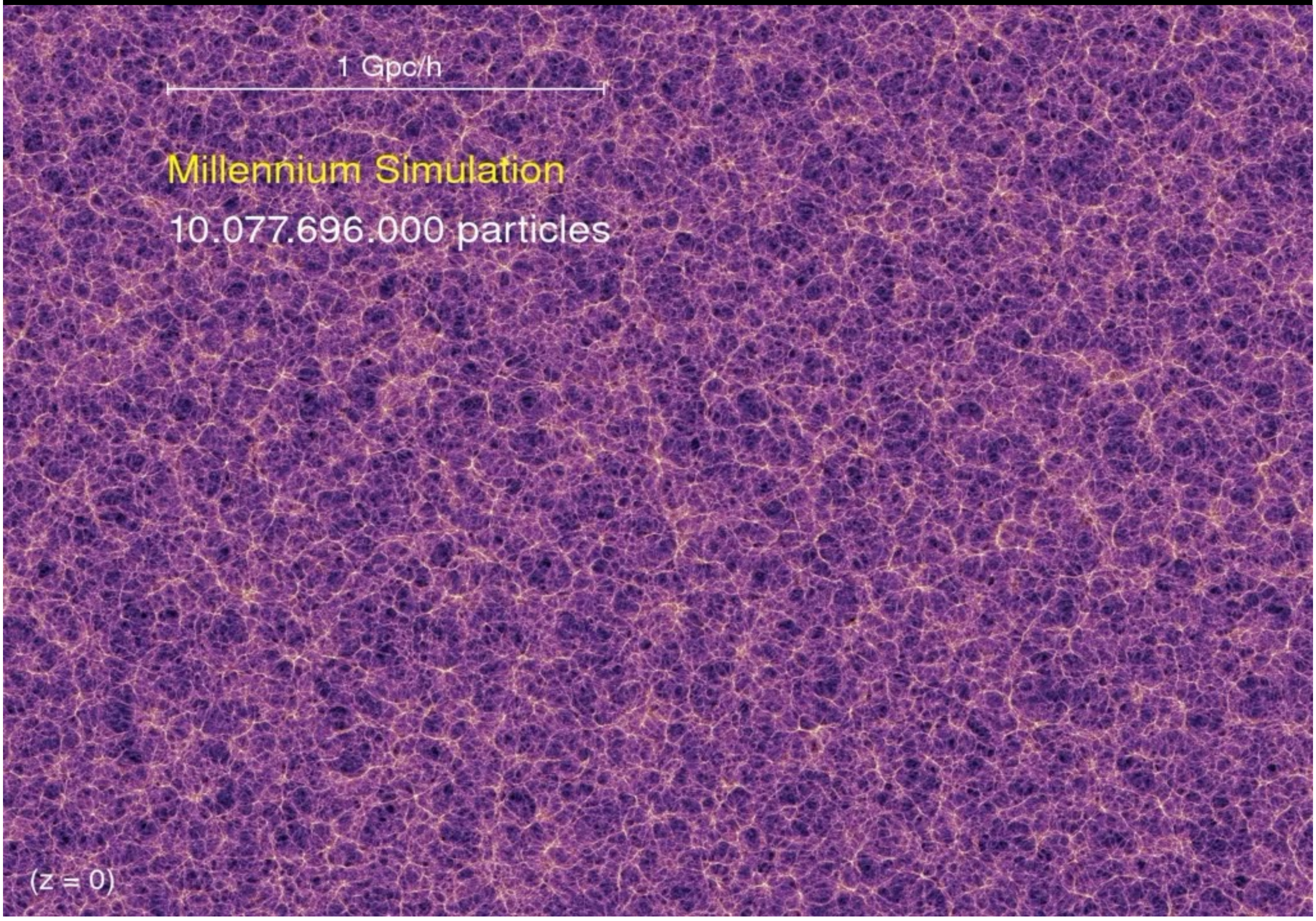

Density fluctuations: t=9 billion years
500 Mpc/h
$$\mathbf{g}(\mathbf{r},t) = -\frac{1}{a}\nabla\phi = \frac{3\Omega H^2}{8\pi}\int \mathrm{d}\mathbf{x}'\,\delta(\mathbf{x}',t)\,\frac{(\mathbf{x}'-\mathbf{x})}{|\mathbf{x}'-\mathbf{x}|^3}$$
Millennium simulation, Springel et al.



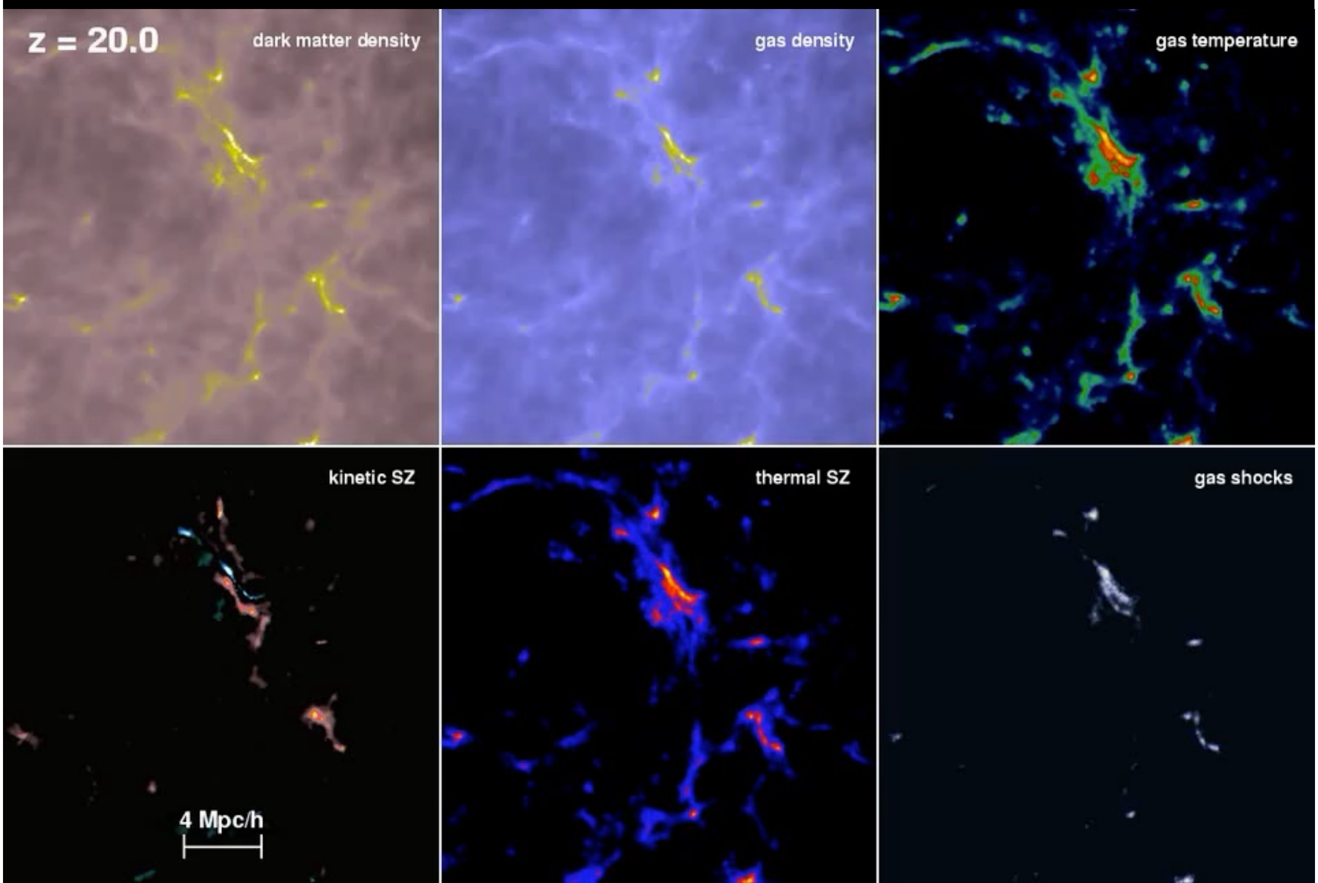
Cosmological structure formation



Large Scale Structure (LSS)



Large Scale Structure (LSS)



The history of the Universe: